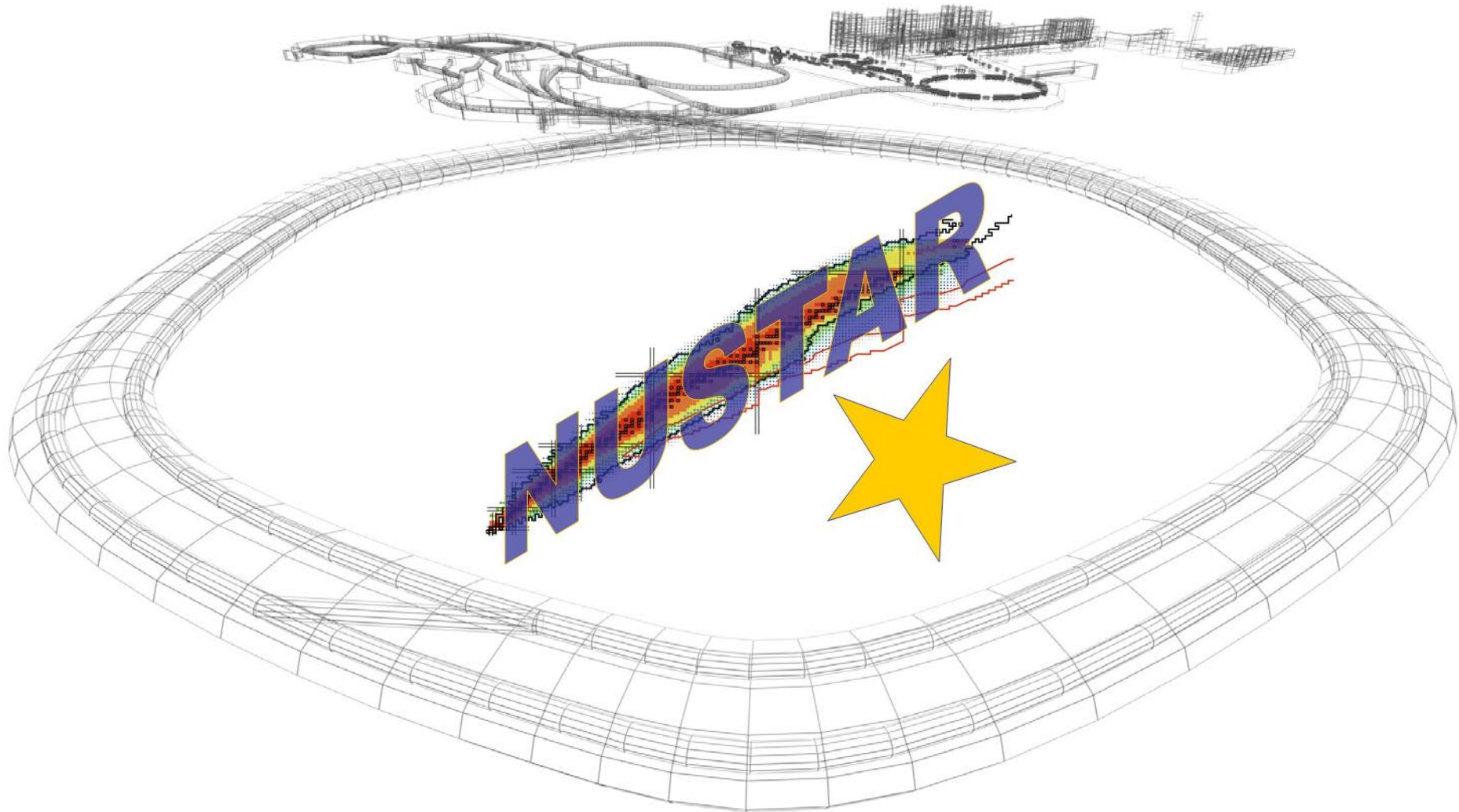


Physics with Exotic Nuclei

Hans-Jürgen Wollersheim



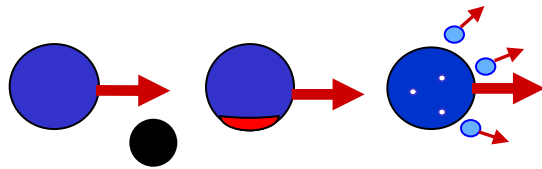
Outline

❖ Excited Fragments – Gateway to Nuclear Structure

- Nuclear Isomers (shape-, spin-, K-traps)
- In-Flight Separation of excited **R**adioactive **I**on **B**eams
- Stopped Beam Experiments and Limitations to Decay Spectroscopy
- Nuclear Shell Closure in ^{98}Cd and ^{132}Cd
- Seniority Isomers in ^{210}Pb , ^{212}Pb , ^{214}Pb , ^{216}Pb
- $T=1$ Isospin Symmetry – Mirror Nuclei
- Silicon Implantation Detector – β -Decay of ^{100}Sn

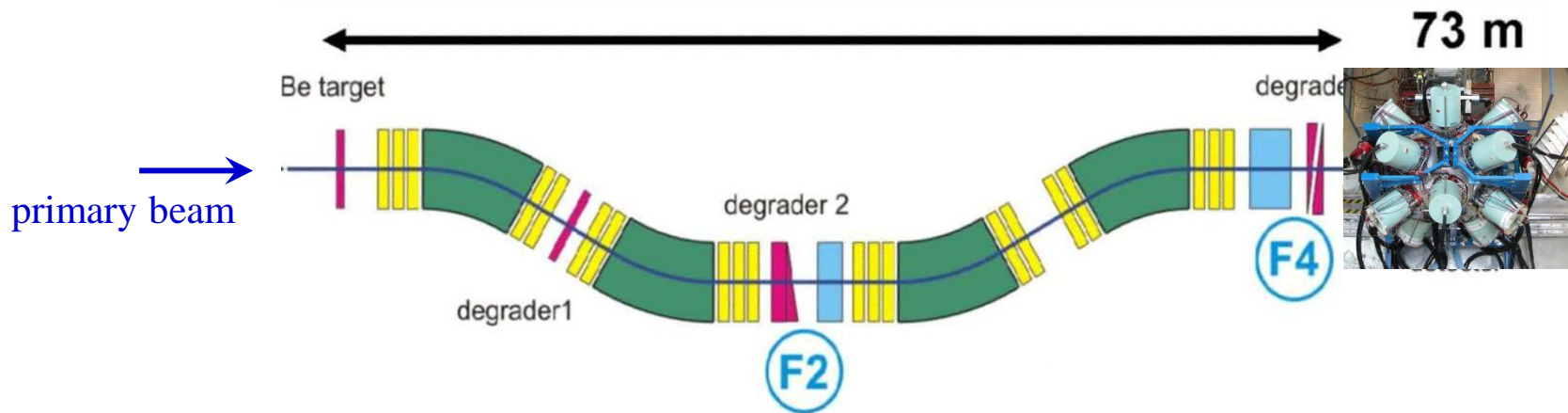
❖ Scattering Experiments with **RIB**s – Nuclear Structure Results

Production of Radioactive Ion Beams



Fragmentation

in ~20% of all cases the fragment is excited



time-of-flight through the fragment separator FRS ~300 ns

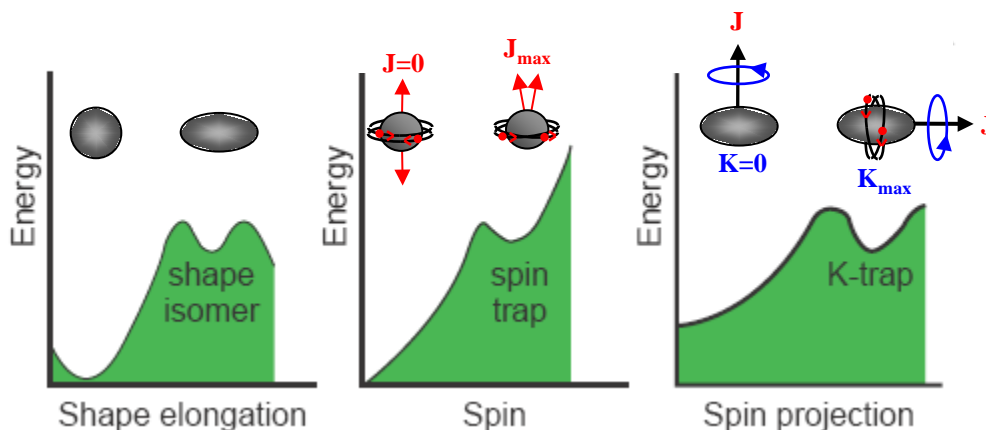
Isomeric states can be investigated!

What is a Nuclear Isomer?

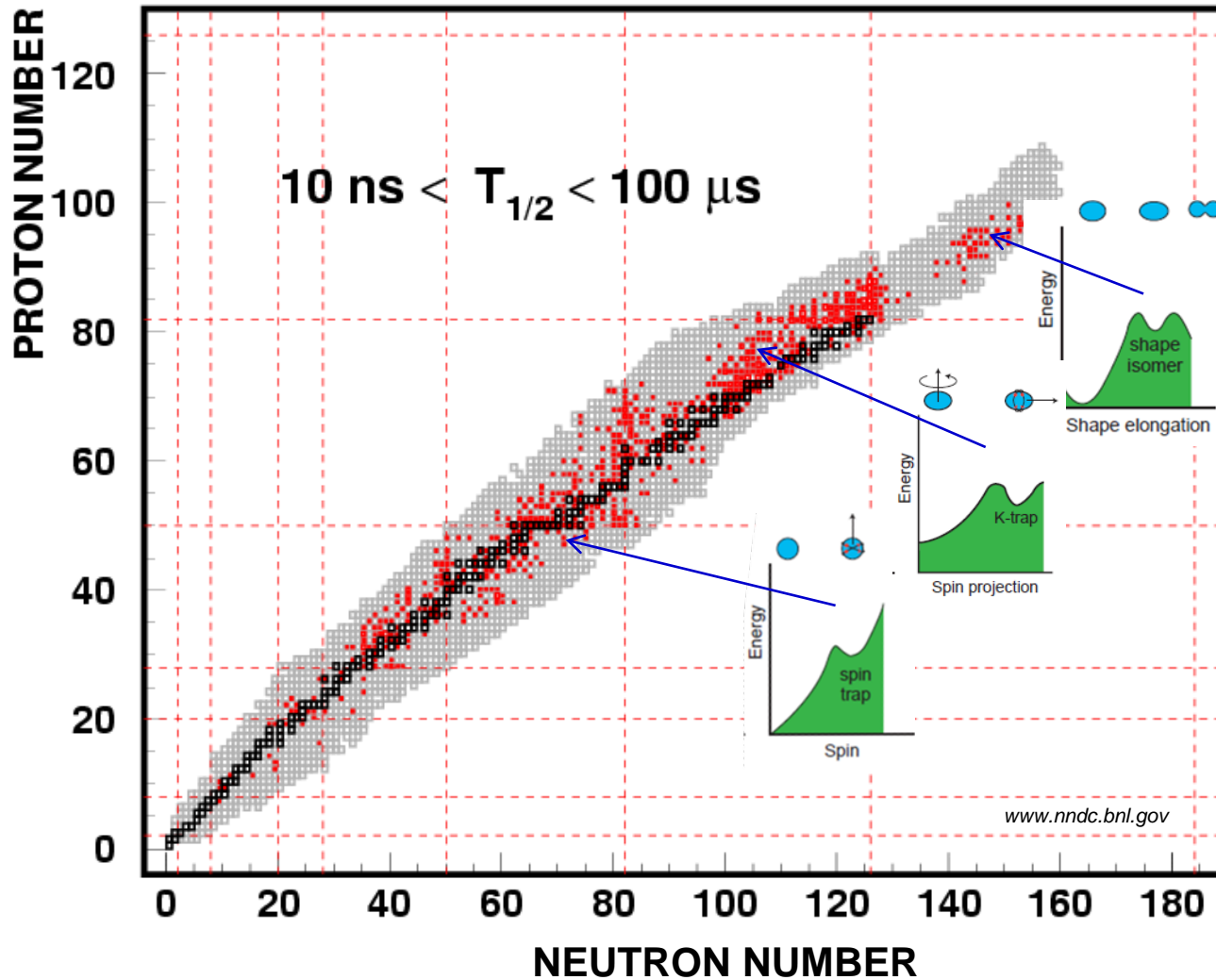
Nuclear Isomer – a long-lived excited nuclear state ($T_{1/2} > 1 \text{ ns}$)
decays by emission of α , β , γ , p , fission, cluster

*The first one discovered by O. Hahn in Berlin in 1921 – decay of ^{234}Pa (70 s)
von Weizsacker, A. Bohr & B. Mottelson*

$$1/\tau \sim E_{\gamma}^{2\lambda+1} |\langle \psi_f | \mathbf{T} | \psi_i \rangle|$$

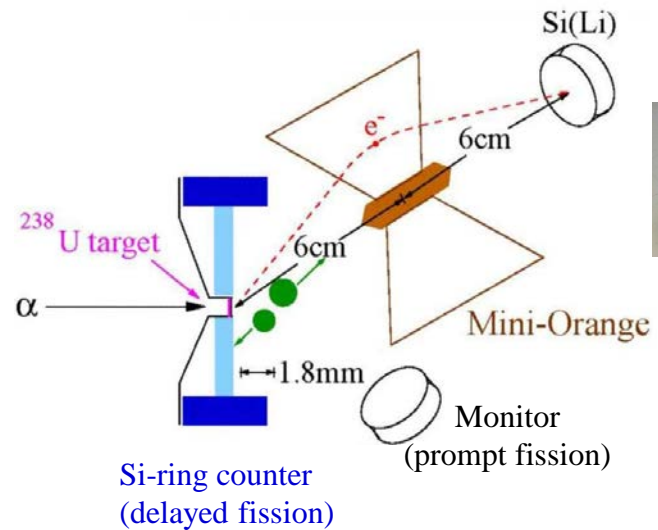
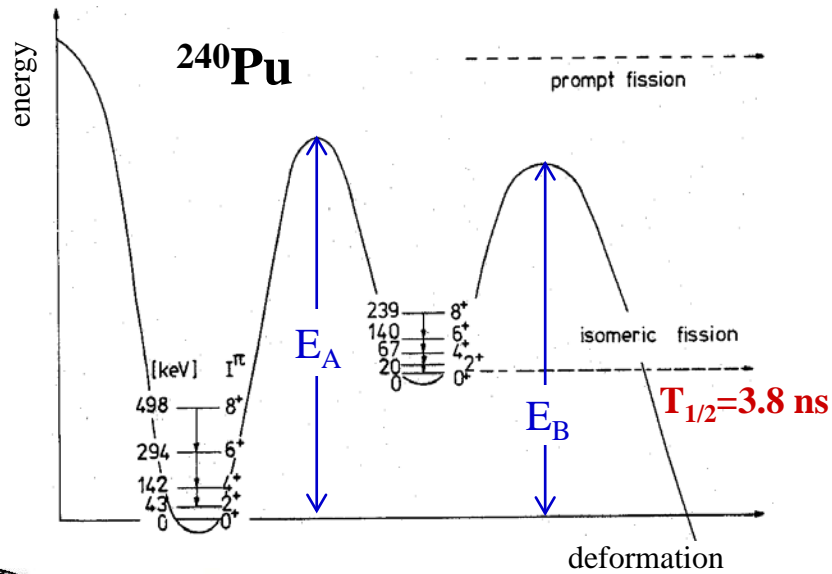


Three Types of Isomers



1. Shape Isomers

- **fission isomer** (discovered by S.M. Polikanov Sov. Phys. JEPT 15 (1962) 105)



$$\frac{\hbar^2}{2\mathcal{I}} = 3.34 \text{ keV}$$

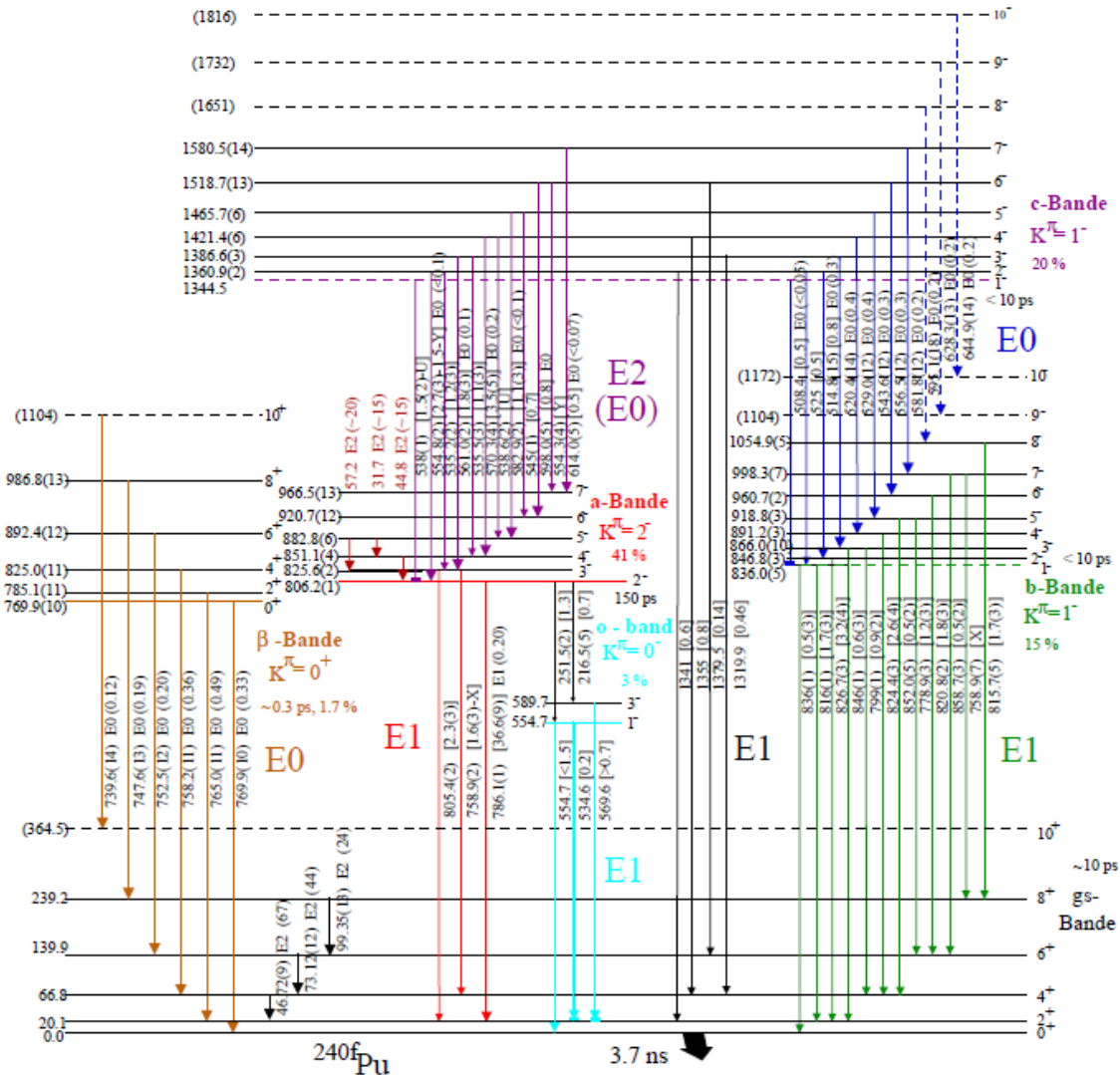
axis ratio 2:1

$$E_A = 5.8 \pm 0.3 \text{ MeV}$$

$$E_B = 5.45 \pm 0.3 \text{ MeV}$$

D. Pansegrau et al., Phys. Lett. B484 (2000) 1
D. Gassmann et al., Phys. Lett. B497 (2001) 181

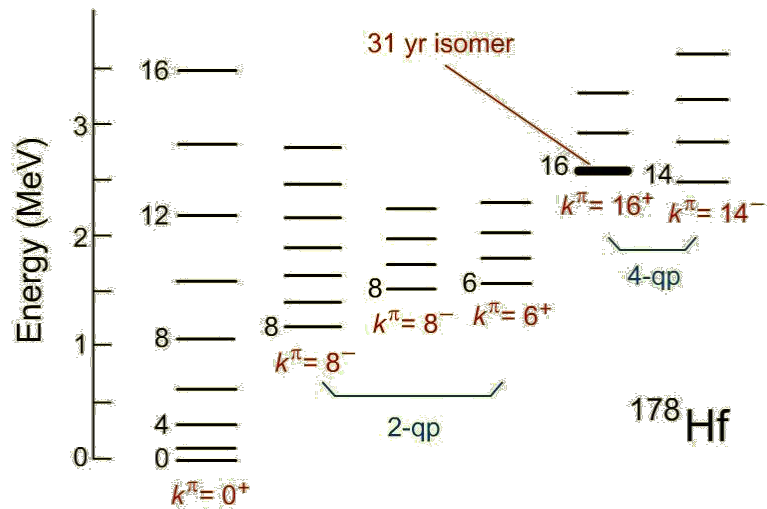
1. Shape Isomers



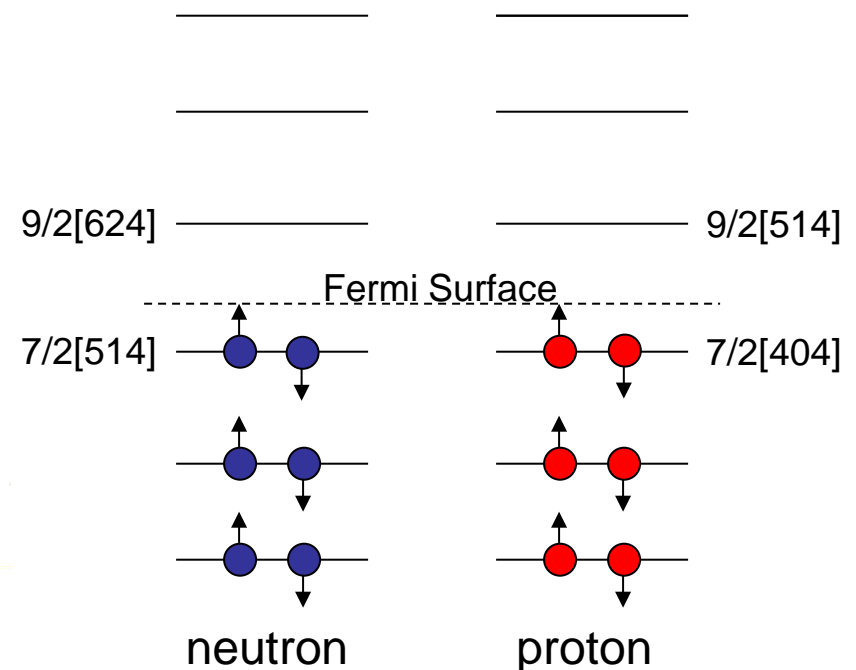
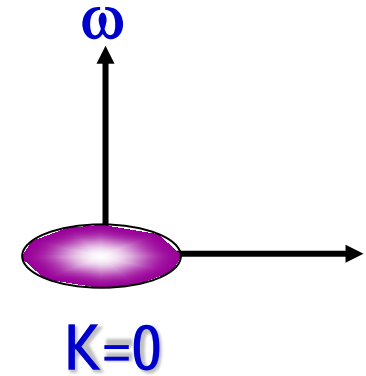
2. K Isomers

- A well-known example:

High-K isomers in ^{178}Hf



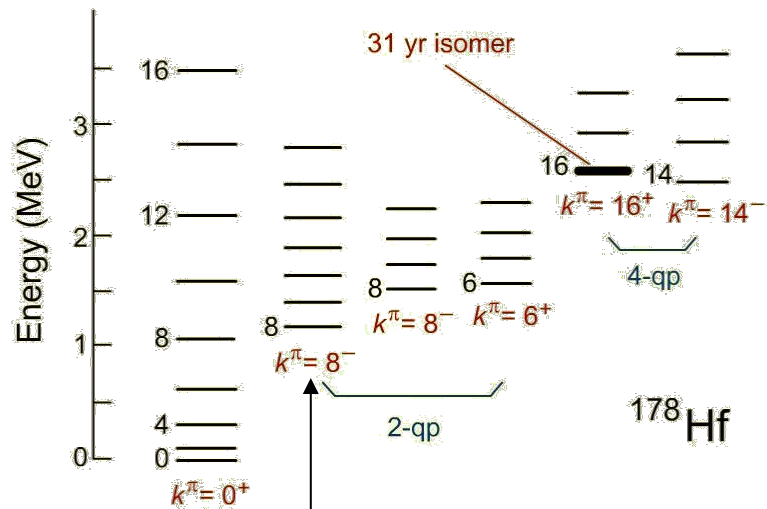
Mullins et al., *Phys. Lett. B* 393 (1997) 279



2. K Isomers

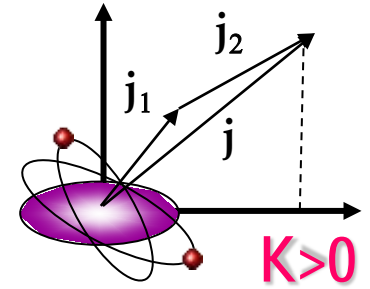
- A well-known example:

High-K isomers in ^{178}Hf

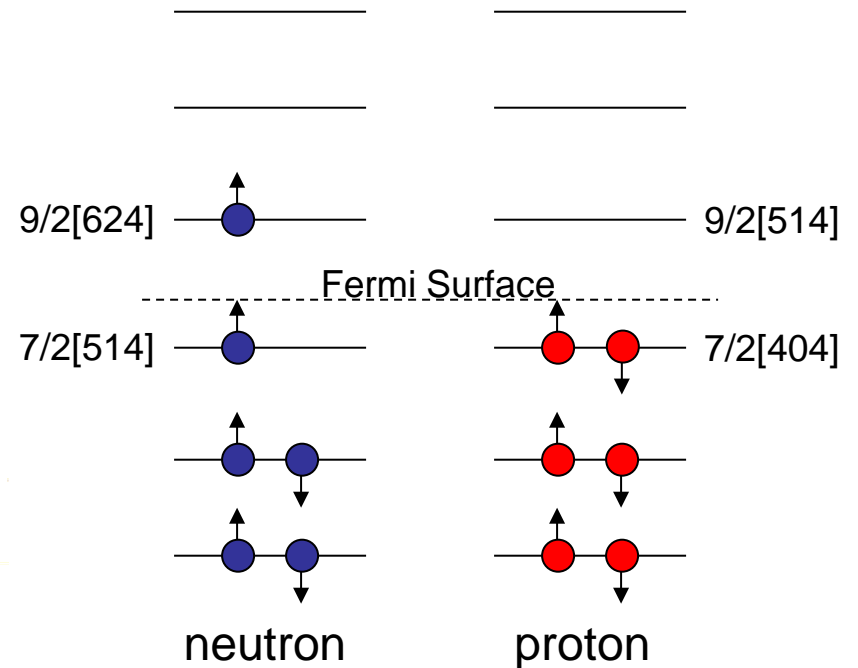


Mullins et al., *Phys. Lett. B* 393 (1997) 279

$$V_{8^-}^2$$

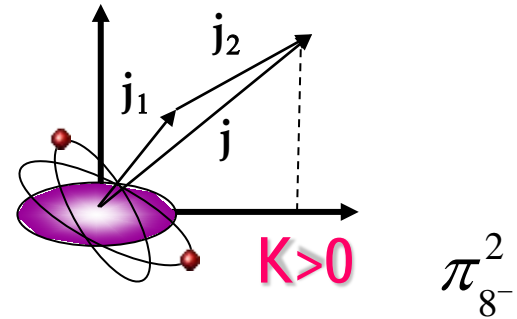


$$V_{8^-}^2$$

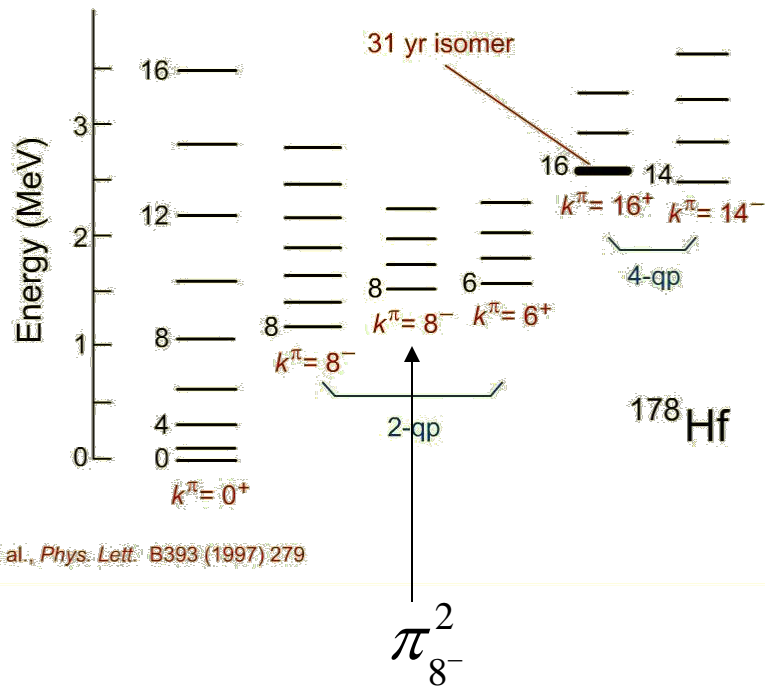


2. K Isomers

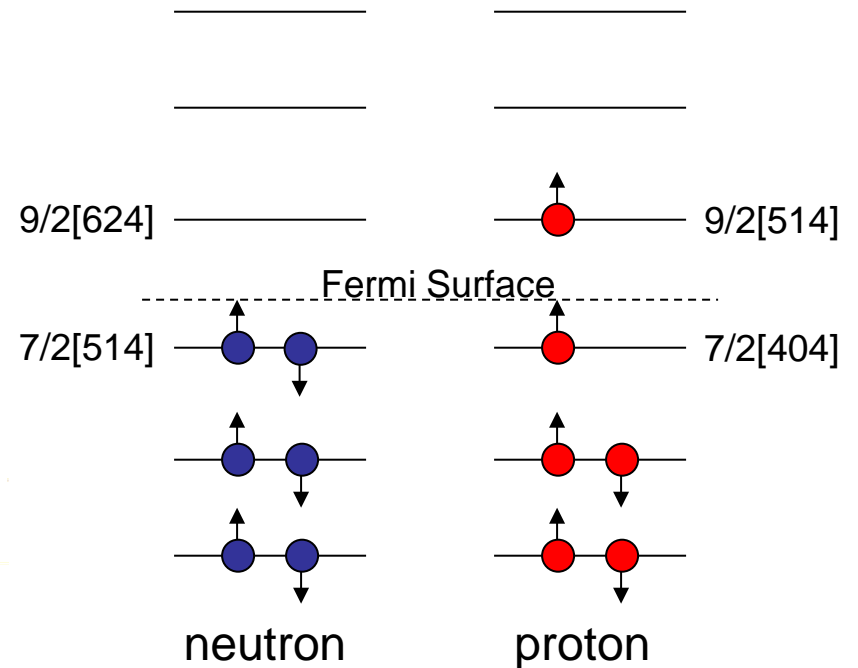
- A well-known example:



High-K isomers in ^{178}Hf



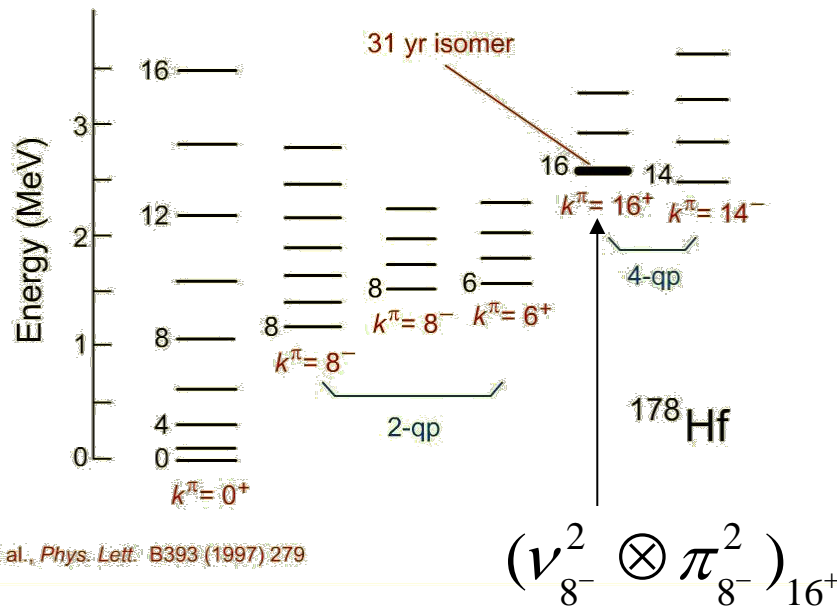
Mullins et al., *Phys. Lett. B* 393 (1997) 279



2. K Isomers

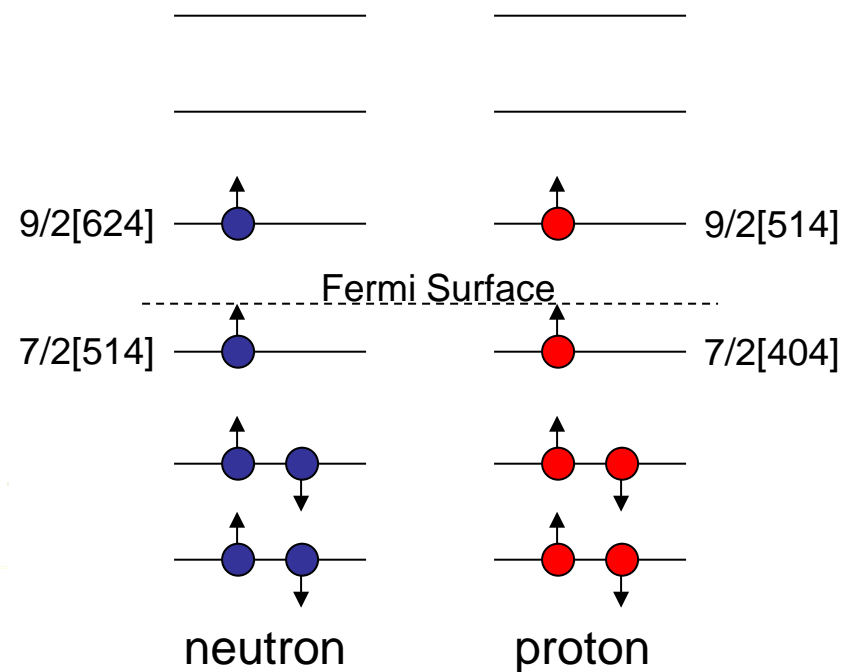
- A well-known example:

High-K isomers in ^{178}Hf



$$(v_{8^-}^2 \otimes \pi_{8^-}^2)_{16^+}$$

$$v_{8^-}^2 \quad \pi_{8^-}^2$$



3. Spin Isomers

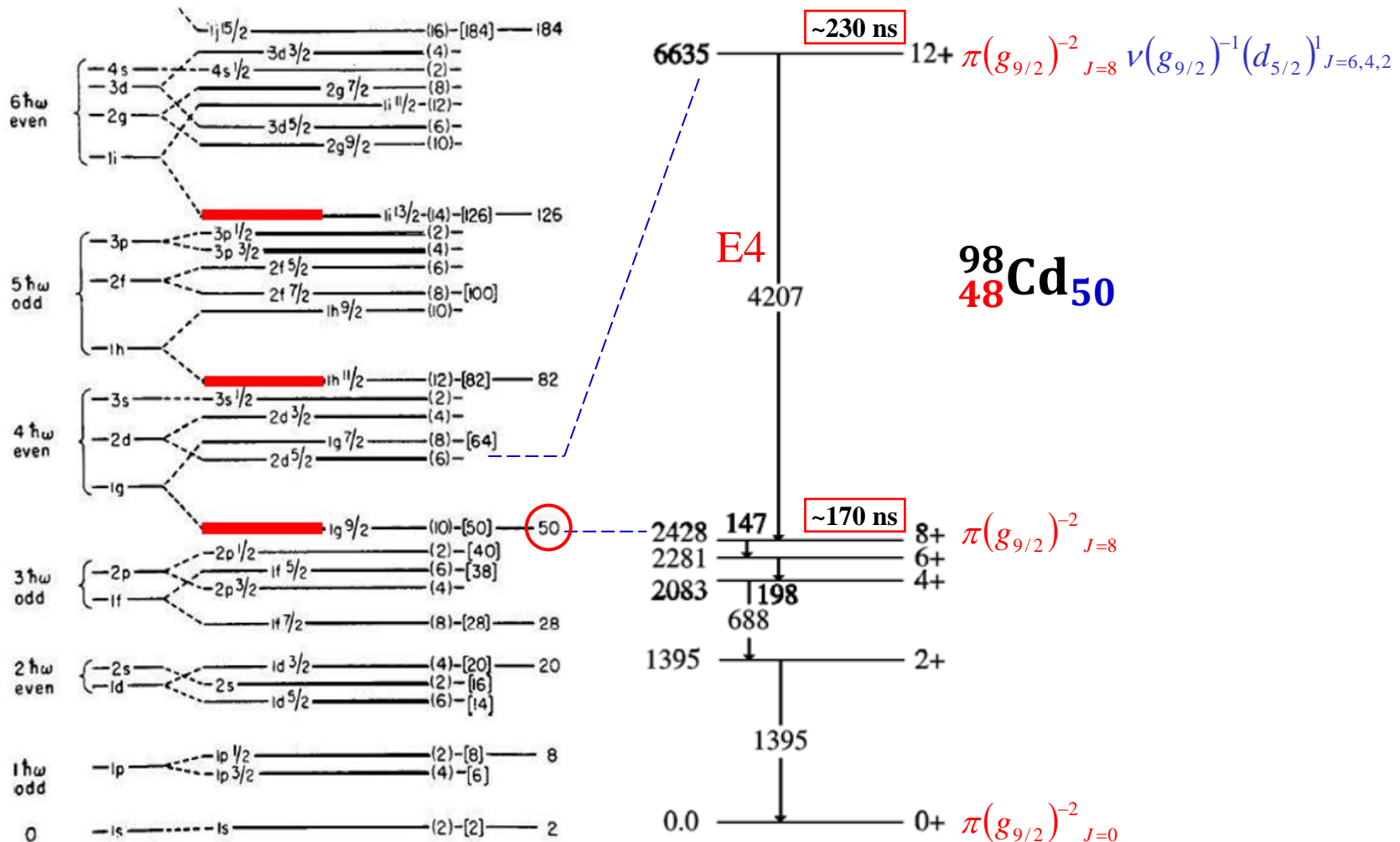
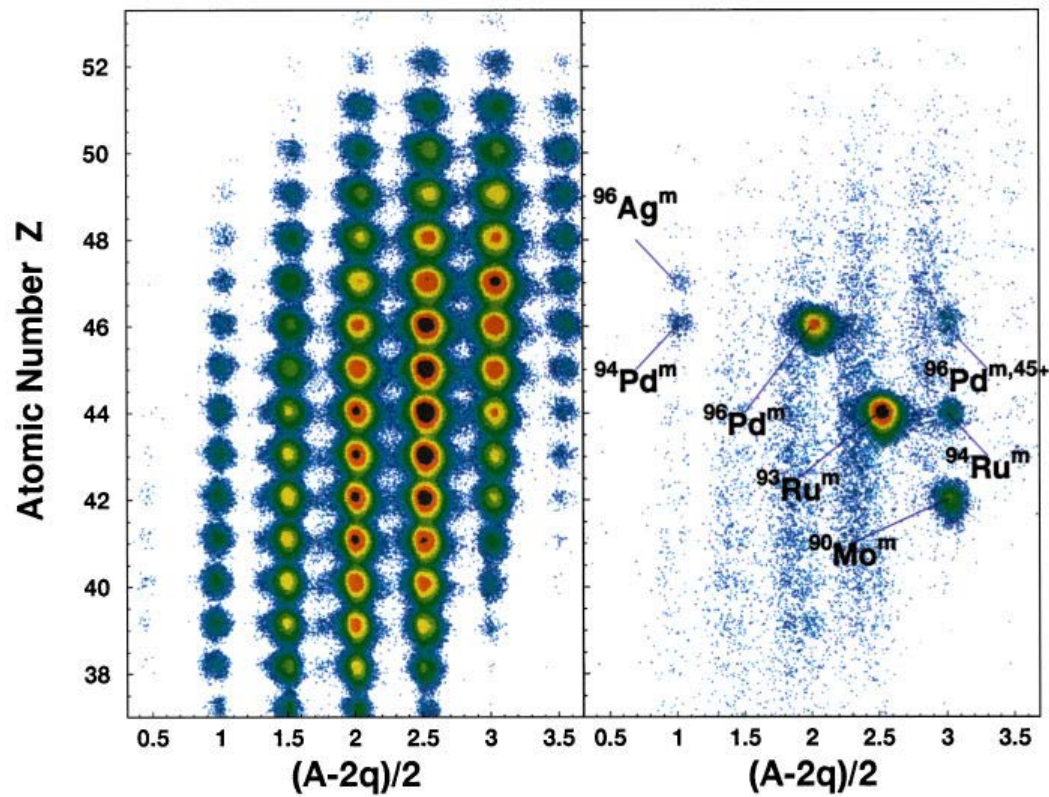
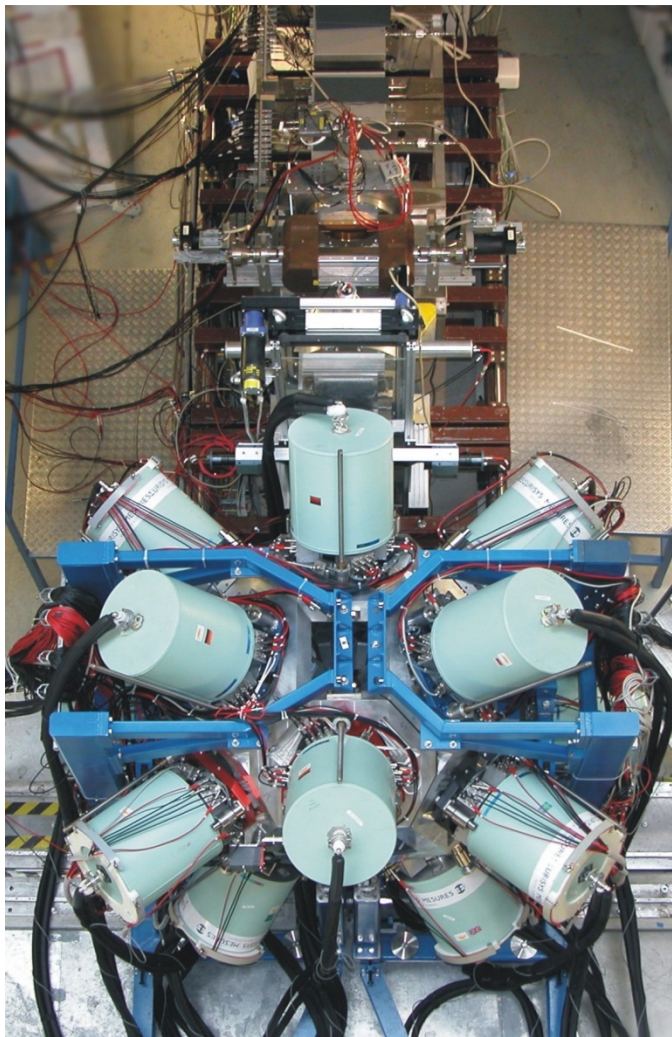
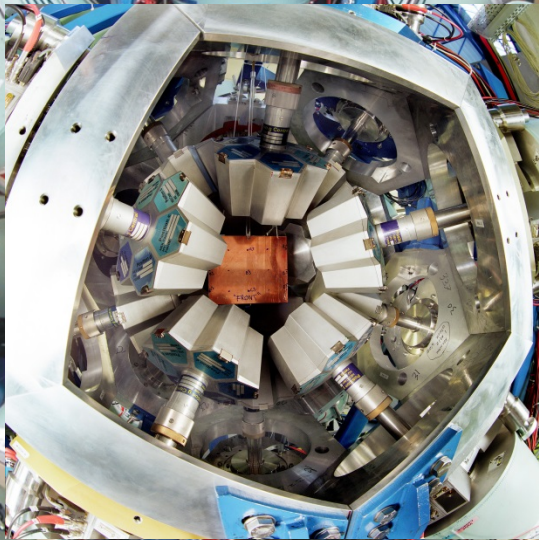


Fig. 7. Realistic level diagram for protons.

Experimental set-up for isomer decay





scintillator
(SC41)

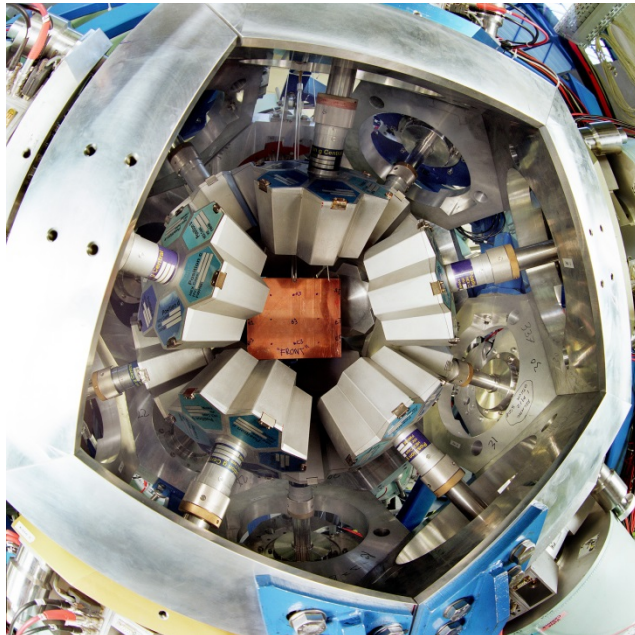
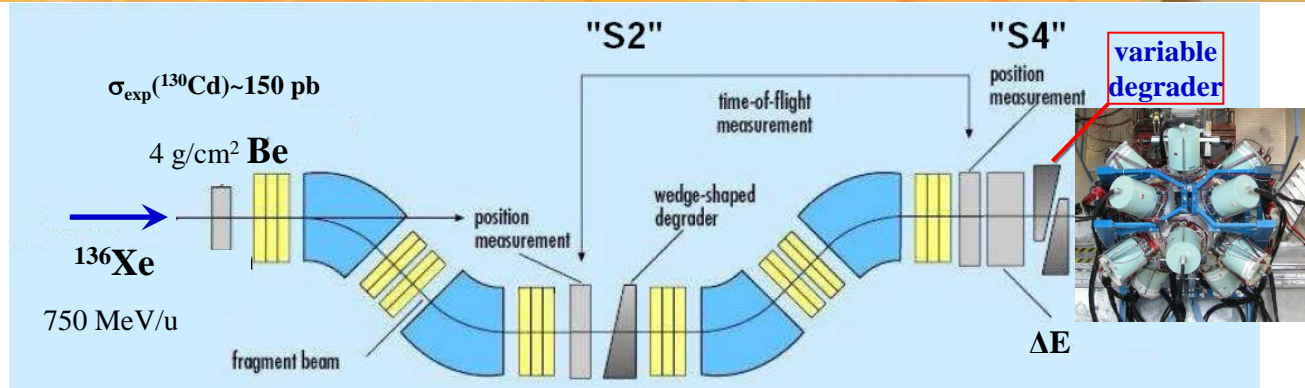
ionization
chambers
(MUSIC41,42)

beam

degrader

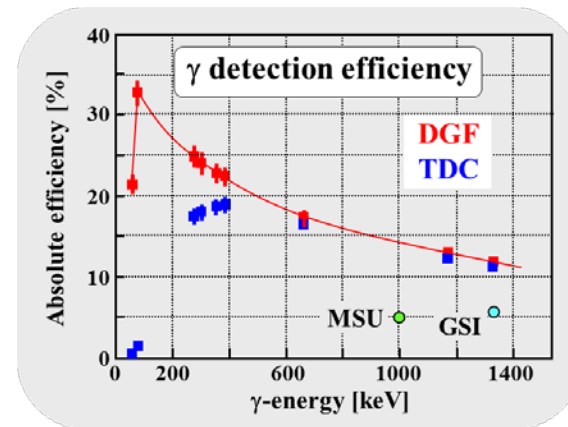
multiwire
chambers
(MW41,MW42)

Experimental set-up with passive target



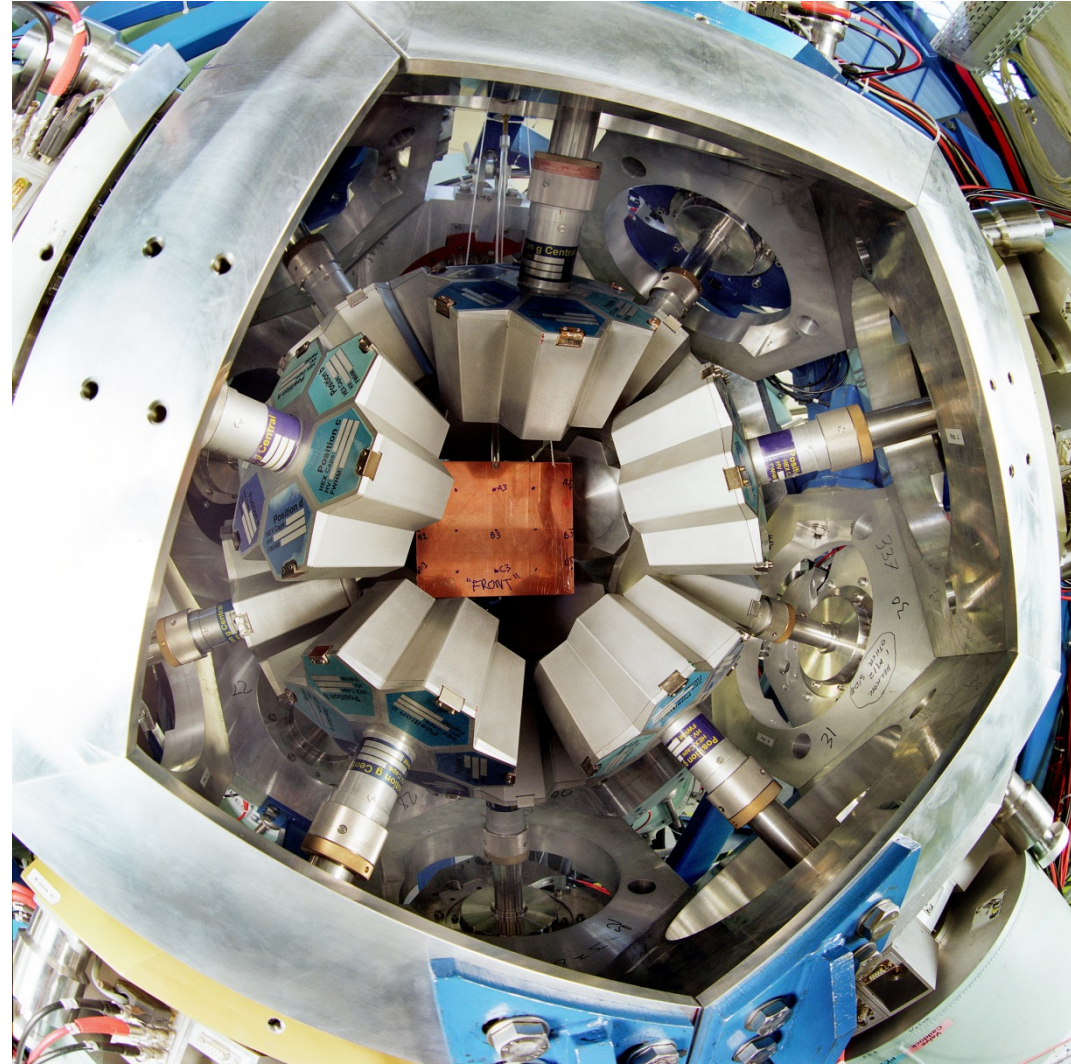
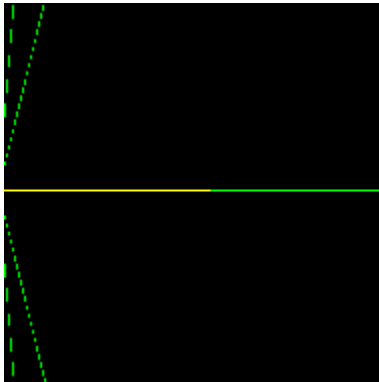
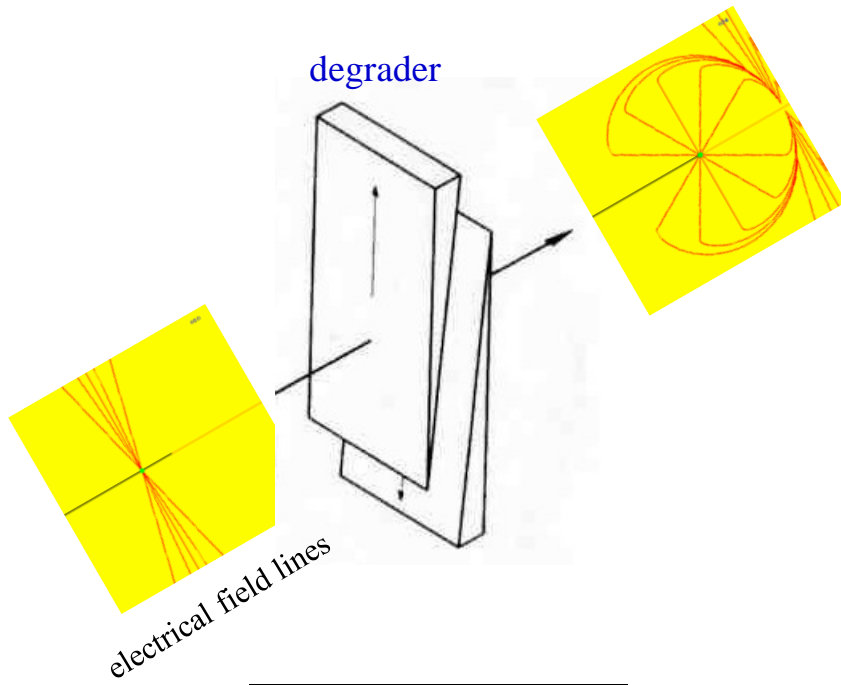
implantation in Cu-plate

15 Cluster detectors with 105 Ge crystals
 $\epsilon_{\gamma} = 11\%$ at 1.3 MeV, 20% at 550 keV, 35% at 100 keV

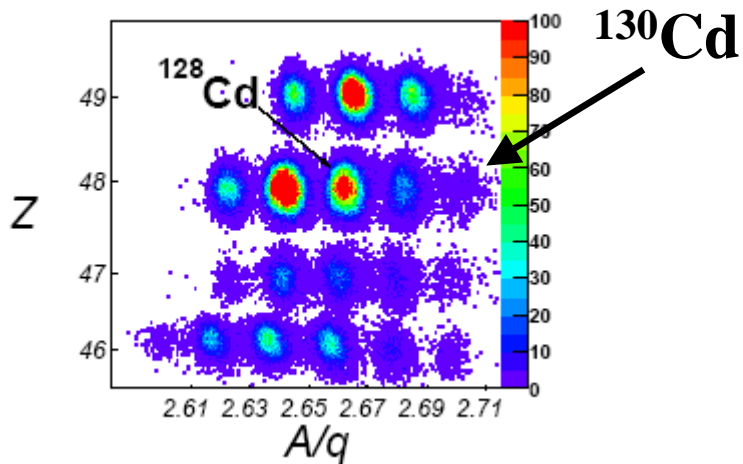
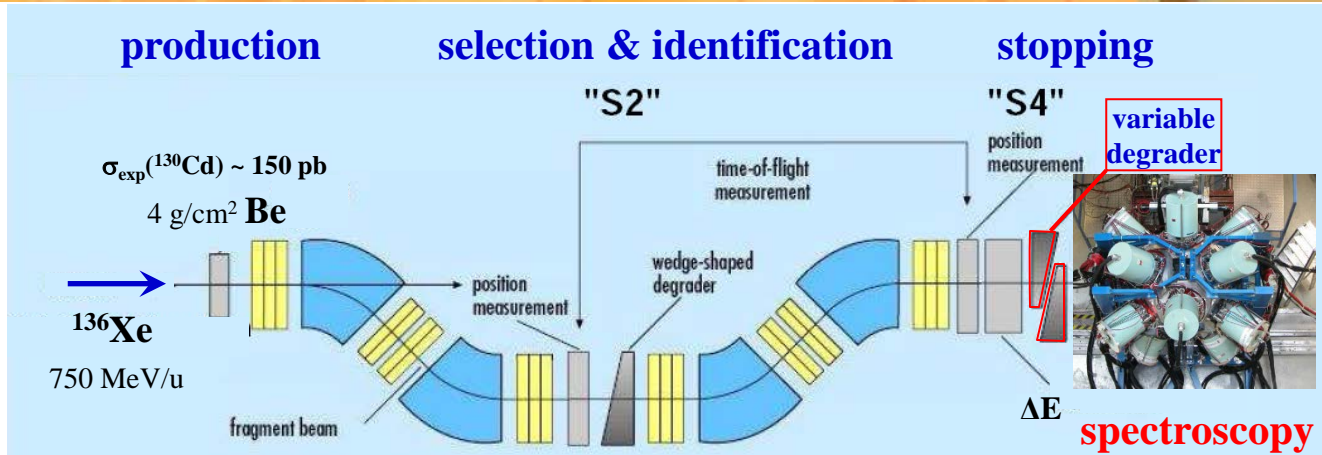


- very high γ -ray efficiency
- high granularity (prompt flash problem)

Limitations to Isomer Spectroscopy



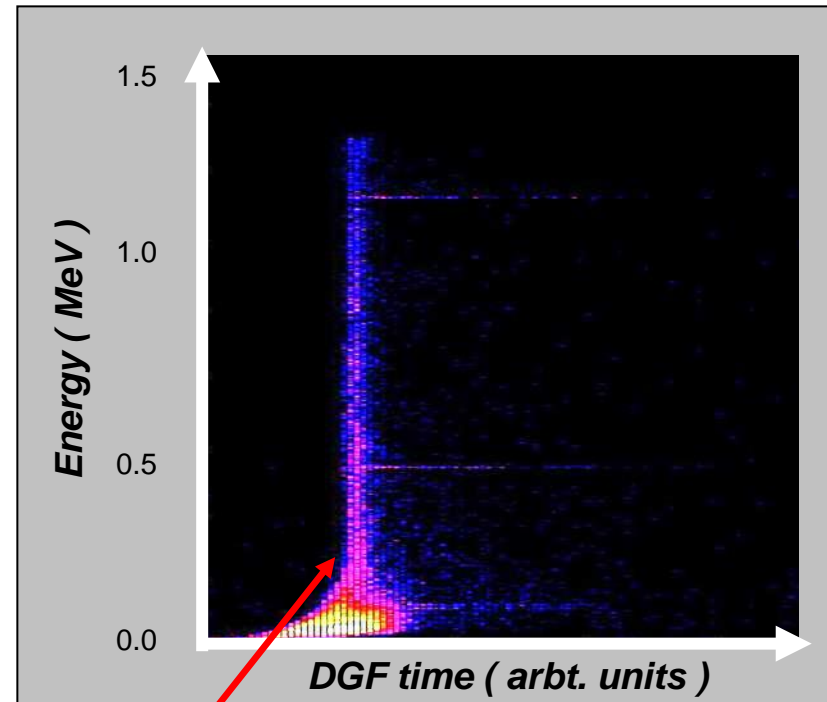
Identification of $^{130}_{48}\text{Cd}_{82}$



➡ 4000 identified ^{130}Cd ions
in fragmentation

Limitations to Isomer Spectroscopy

^{130}Cd : **DGF-timing**

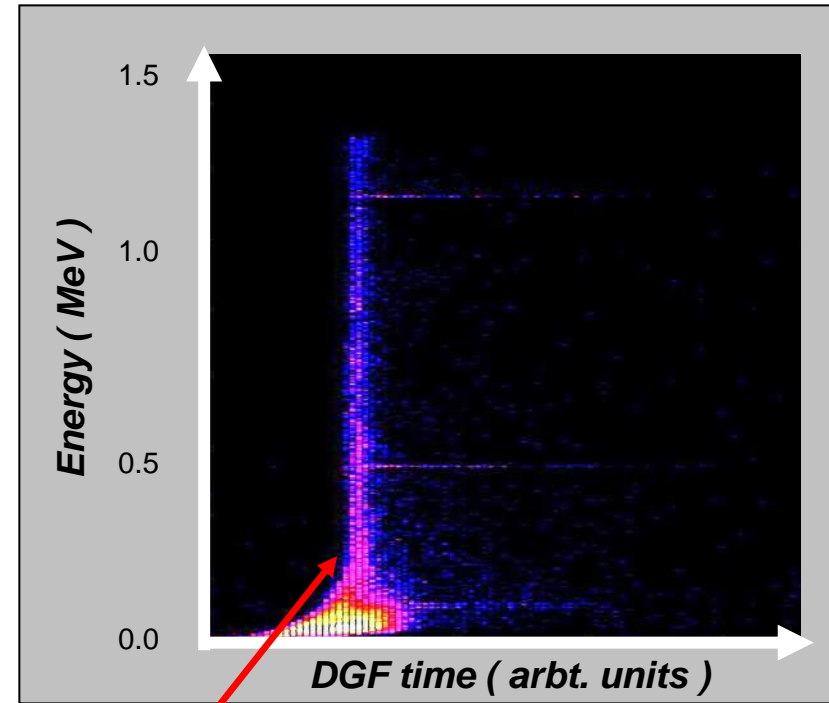
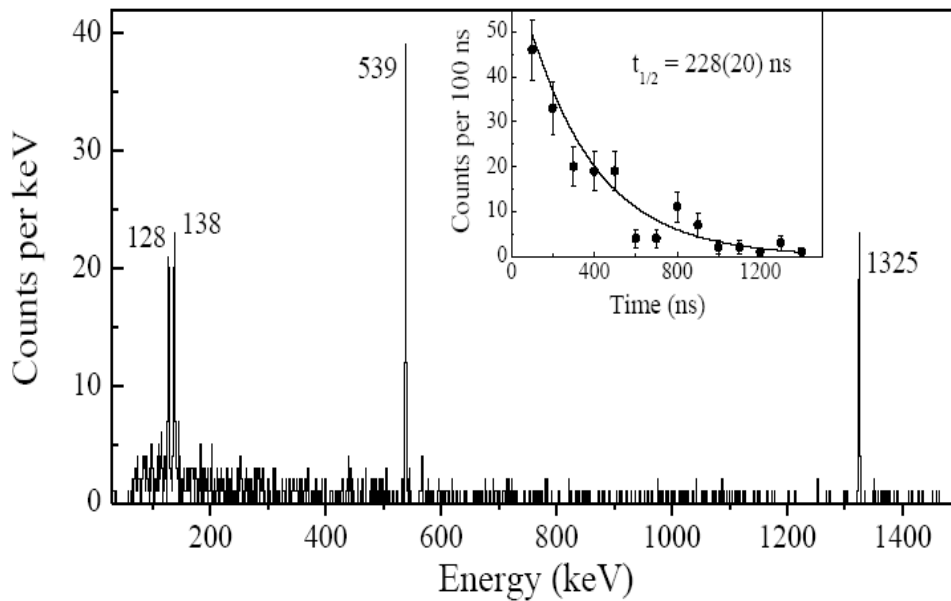


Prompt γ -flash

Decay time range: 20 ns ... 20 μ s

Decay Spectroscopy Probes Shell Closures

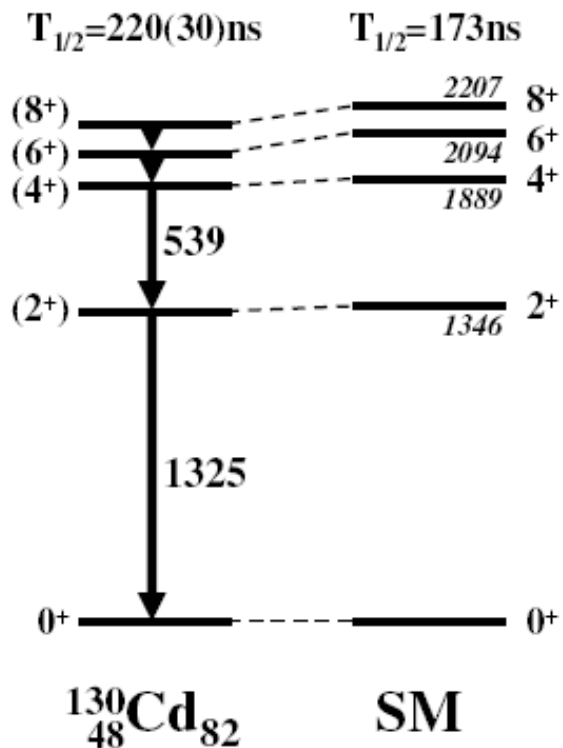
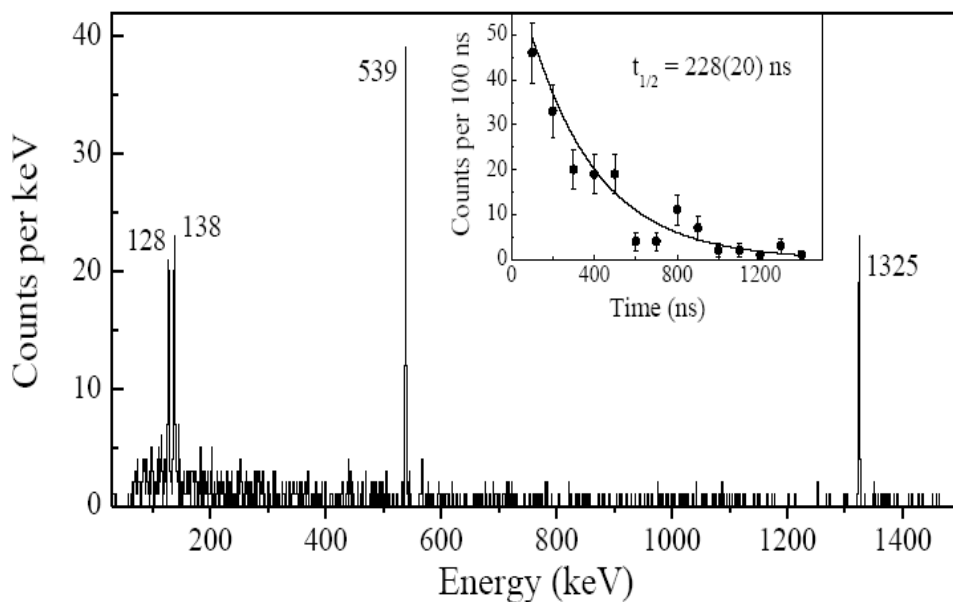
^{130}Cd : **DGF-timing**



Prompt γ -flash

Decay time range: 20 ns ... 20 μ s

Decay Spectroscopy Probes Shell Closures



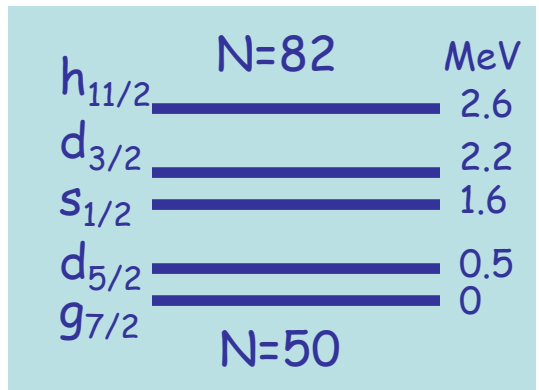
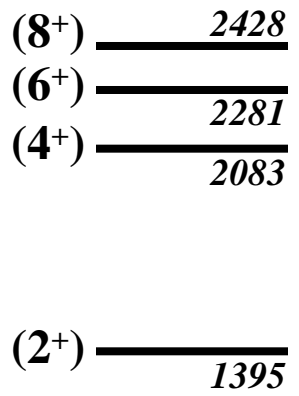
No Shell quenching observed

$8^+(g_{9/2})^{-2}$ Seniority Isomers in ^{98}Cd and ^{130}Cd

Sn100 0.94 s 0+	Sn101 3 s 0+	Sn102 4.5 s 0+	Sn103 7 s 0+	Sn104 20.9 s 0+	Sn105 31 s 0+	Sn106 115 s 0+	Sn107 2.90 m (5/2+)	Sn108 10.50 m 0+	Sn109 18.0 m 5/2(+)	Sn110 411 h 0+	Sn111 35.3 m 7/2+	Sn112 0+	Sn113 115.09 d 1/2+	Sn114 0+	Sn115 1/2+	Sn116 0+	Sn117 1/2+	Sn118 0+	Sn119 3/2+	Sn120 0+	Sn121 17.04 h 3/2+	Sn122 0+	Sn123 119.2 d 11/2-	Sn124 0+	Sn125 9.64 d 11/2-	Sn126 1E+5 y 0+	Sn127 2.10 h (11/2-)	Sn128 59.07 m 0+	Sn129 2.13 m (3/2+)	Sn130 3.72 m 0+	Sn131 56.9 s (3/2+)	Sn132 39.7 s 0+	
In99 7.0 s 0+	In100 15.1 s 0+	In101 22 s (6+)	In102 1.30 m (9/2+)	In103 5.97 m (6+)	In104 32.4 m (9/2+)	In105 6.2 m 7-	In106 32.4 m 9/2+	In107 4.9 h 9/2+	In108 2.8047 d 9/2+	In109 4.2 h 7+	In110 4.9 h 9/2+	In111 2.8047 d 9/2+	In112 14.97 m 1+	In113 71.9 s 1+	In114 4.41E+14 y 9/2+	In115 7.08 s 9/2+	In116 14.10 s 1+	In117 14.10 s 9/2+	In118 5.0 s 1+	In119 2.4 m 9/2+	In120 3.80 s 1+	In121 23.1 s 9/2+	In122 1.5 s 1+	In123 5.98 s 9/2+	In124 3.11 s 3+	In125 2.36 s 9/2(+)	In126 1.60 s 3(+)	In127 1.00 s (9/2-)	In128 0.84 s (3+)	In129 0.61 s (9/2-)	In130 0.32 s 3(+)	In131 0.282 s (9/2-)	In132 0.282 s (9/2-)
Cd98 9.2 s 0+	Cd99 0+	Cd100 0+	Cd101 1.35 m (5/2+)	Cd102 2.5 m 0+	Cd103 7.1 m (5/2+)	Cd104 57.7 m 0+	Cd105 5.5 m 5/2+	Cd106 1.5 s 0+	Cd107 4.50 h 5/2+	Cd108 462.7 d 0+	Cd109 462.7 d 5/2+	Cd110 12.49 s 0+	Cd111 12.89 s 1/2+	Cd112 24.13 s 0+	Cd113 27.257 y 1/2+	Cd114 28.73 s 0+	Cd115 28.46 h 1/2+	Cd116 7.49 s 0+	Cd117 1.94 h 1/2+	Cd118 1.99 m 0+	Cd119 50.31 s 3/2+	Cd120 50.31 s 0+	Cd121 13.5 s (3/2-)	Cd122 2.34 s 0+	Cd123 2.19 s (3/2+)	Cd124 1.55 s 0+	Cd125 0.65 s (3/2+)	Cd126 1.596 s 0+	Cd127 0.97 s (3/2-)	Cd128 0.24 s 0+	Cd129 0.24 s (3/2-)	Cd130 0.24 s 0+	

Cd98
9.2 s
0+
EC

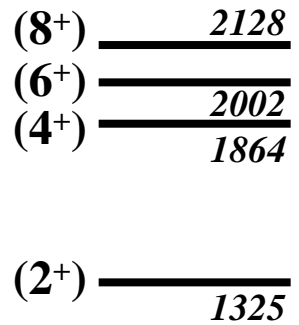
N=50
Z=48



participating neutron-orbitals

Cd130
0.20 s
0+
β-n

N=82
Z=48

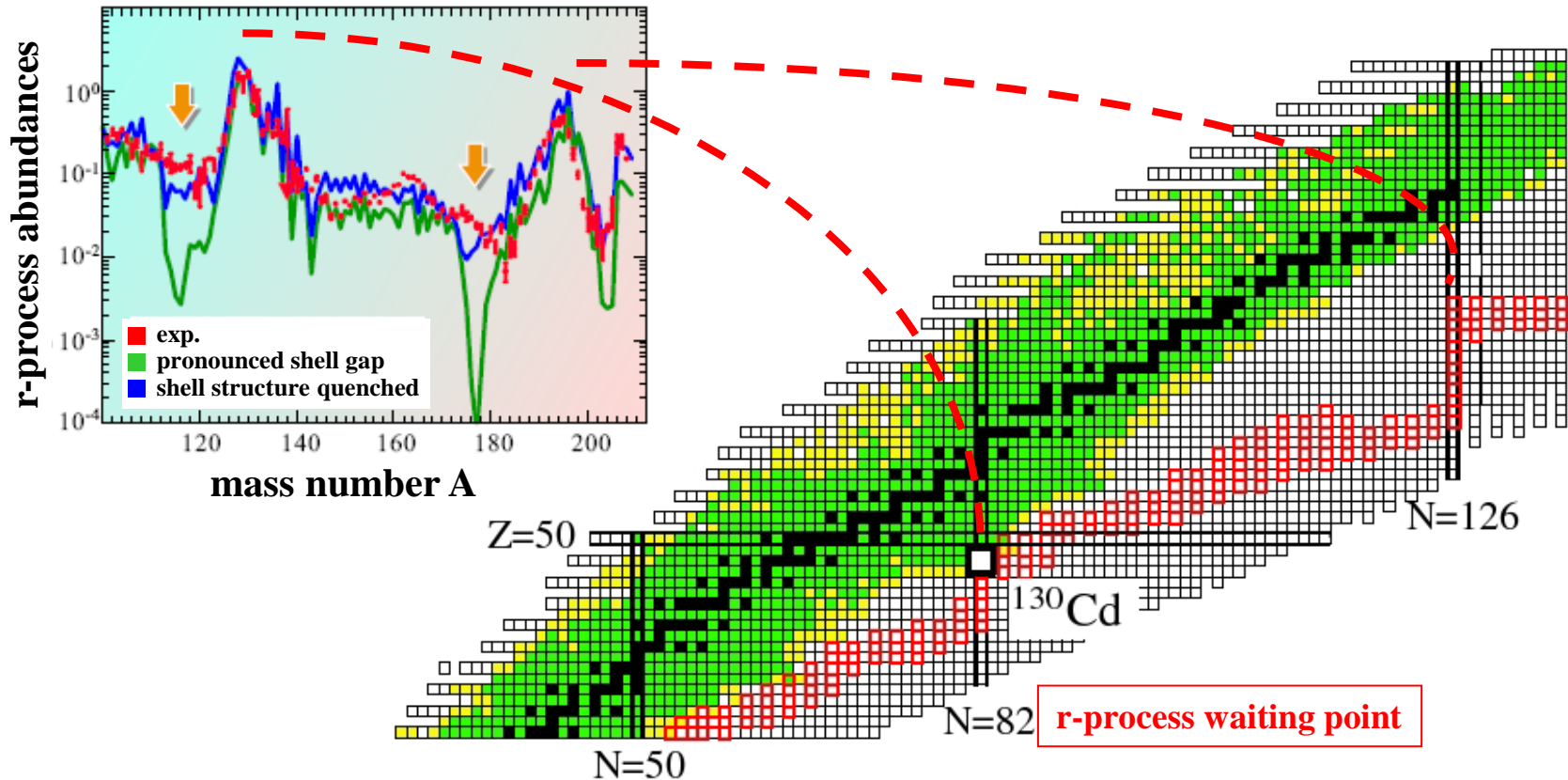


two proton holes in the $g_{9/2}$ orbit

No dramatic shell quenching!

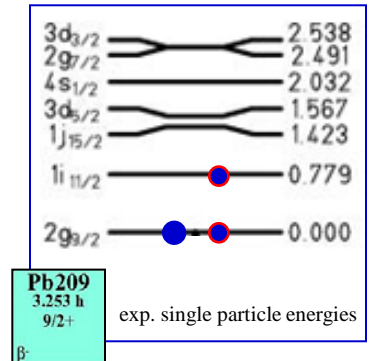
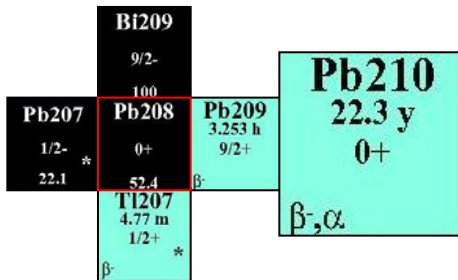


The astrophysical r-process 'path'



Assumption of a N=82 shell quenching leads to a considerable improvement in the global abundance fit in r-process calculations !

Level Scheme of ^{210}Pb



(pairing energy)
residual interaction !

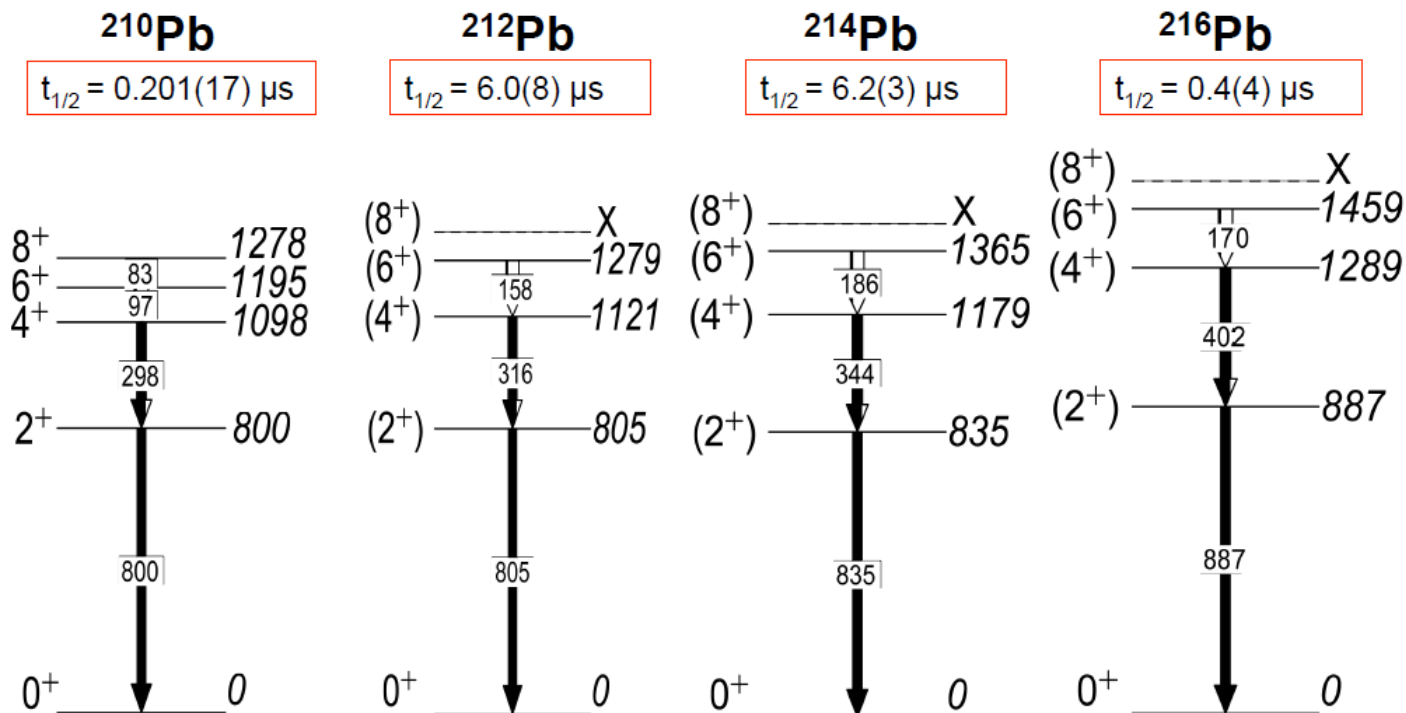


Level Schemes in Neutron-Rich Pb Isotopes

Pb205 1.53E+7 y 5/2- *	Pb206 0+	Pb207 1/2- *	Pb208 0+	Pb209 3.253 h 9/2+	Pb210 22.3 y 0+	Pb211 36.1 m 9/2+	Pb212 10.64 h 0+	Pb213 10.2 m (9/2+)	Pb214 26.8 m 0+	Pb215 36 s (5/2+)	Pb216	Pb217	Pb218
EC	24.1	22.1	52.4	β^-	β^-, α	β^-	β^-	β^-	β^-	β^-			

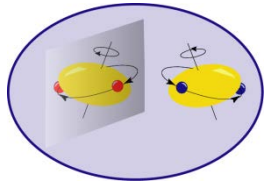


$g_{9/2}$

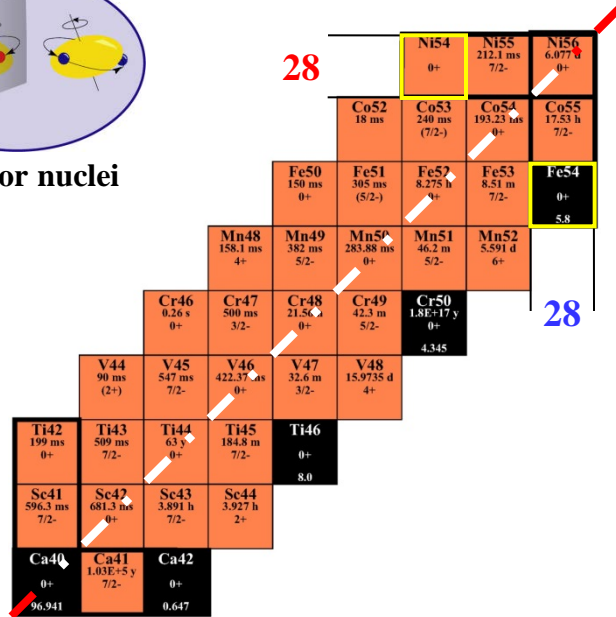


T=1 Isospin Symmetry in pf-shell Nuclei

search for isospin breaking effects

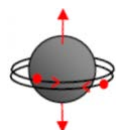
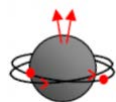
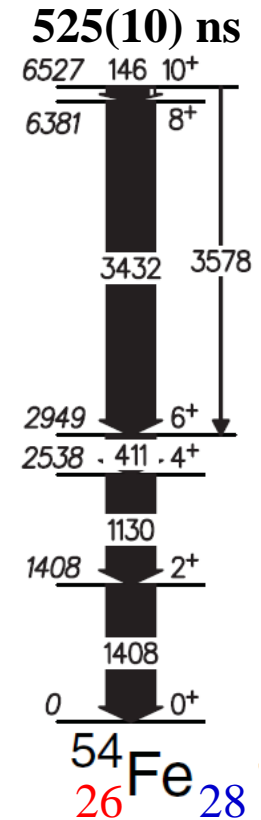
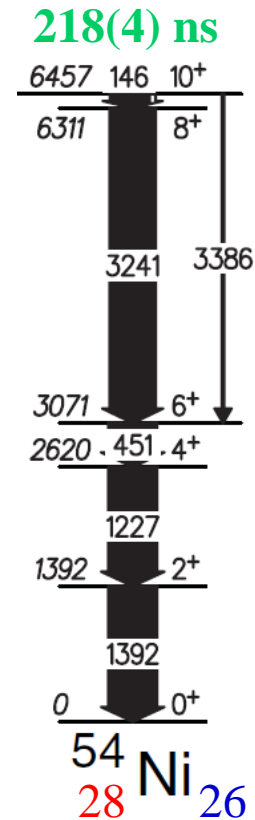


mirror nuclei

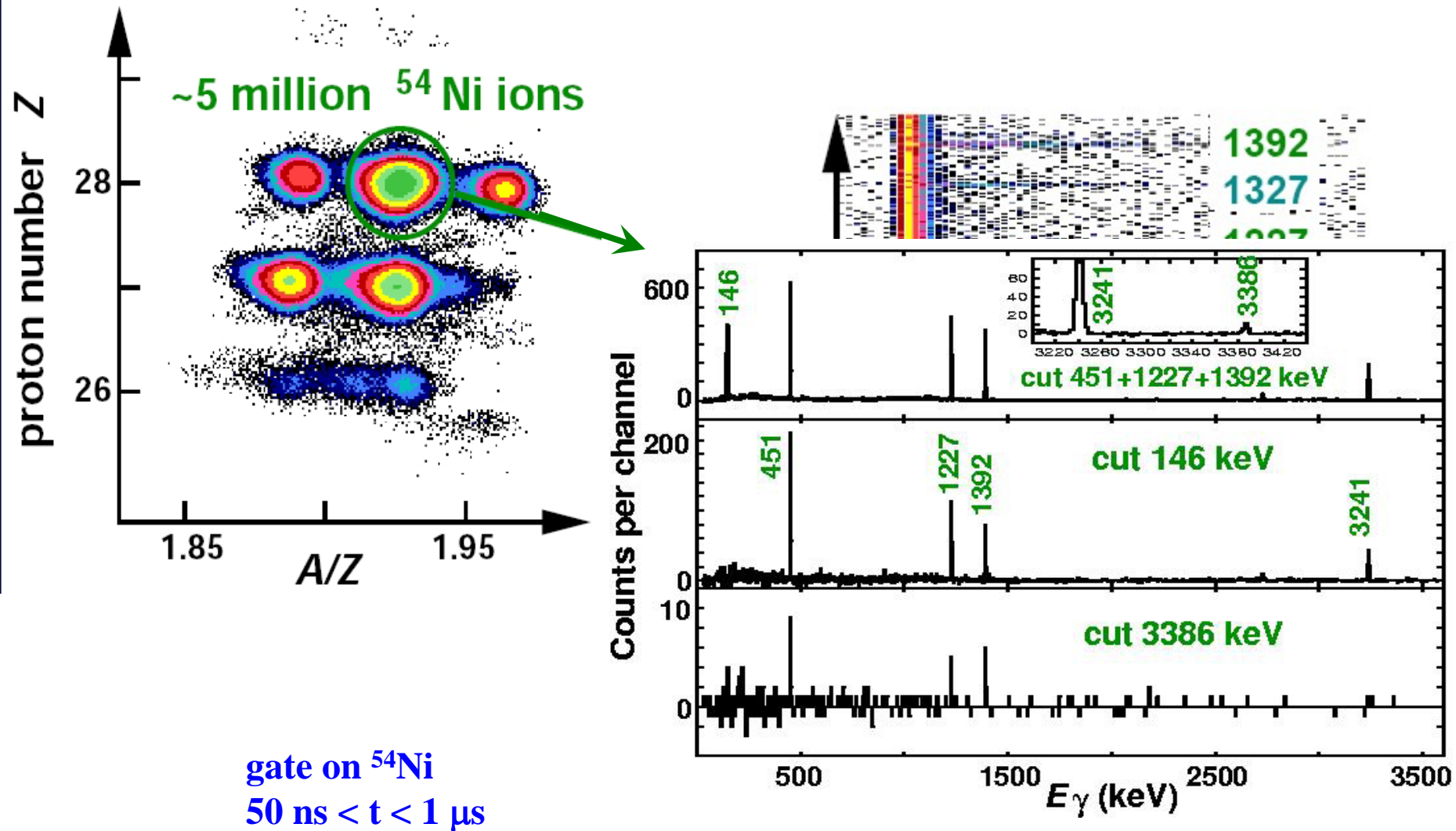


N=Z

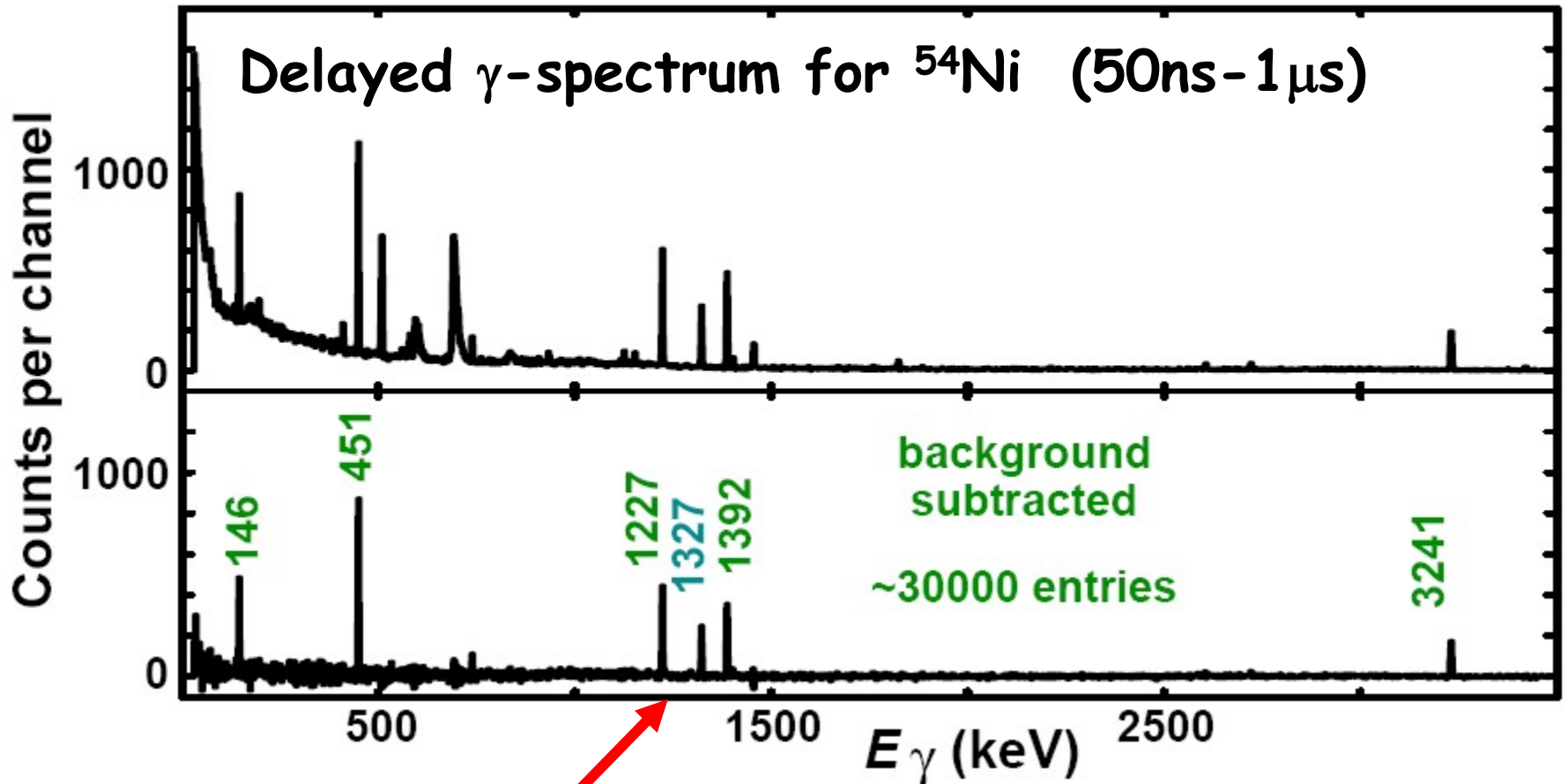
decay of the excited 10⁺-state by proton emission and γ -radiation



Identification of ^{54}Ni



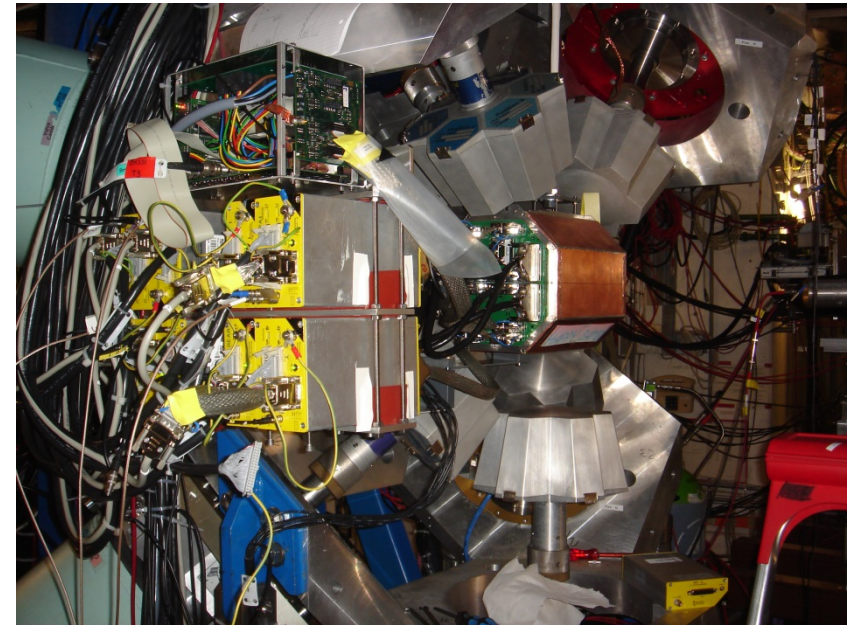
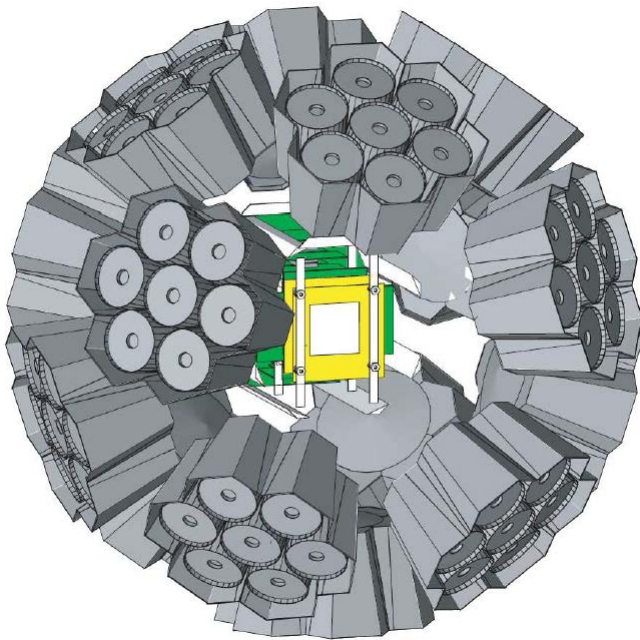
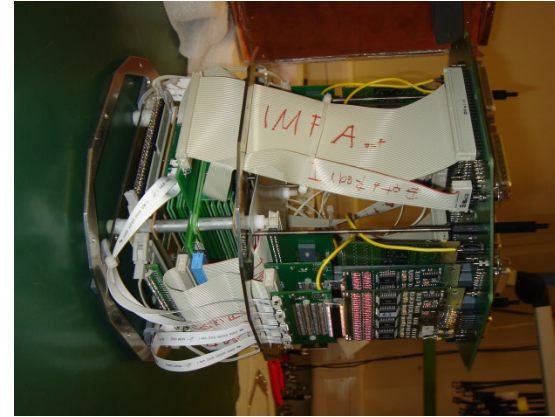
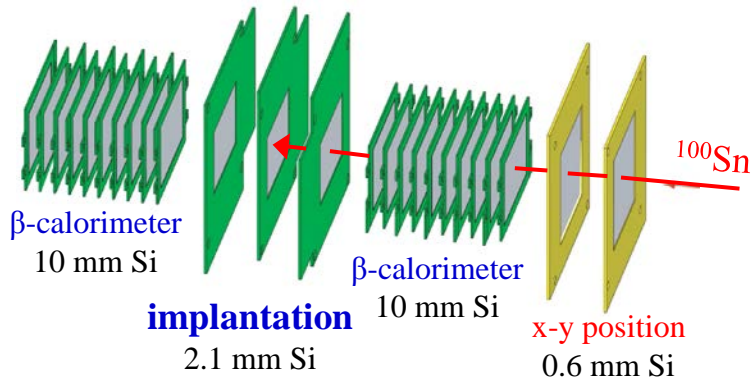
The big surprise ...



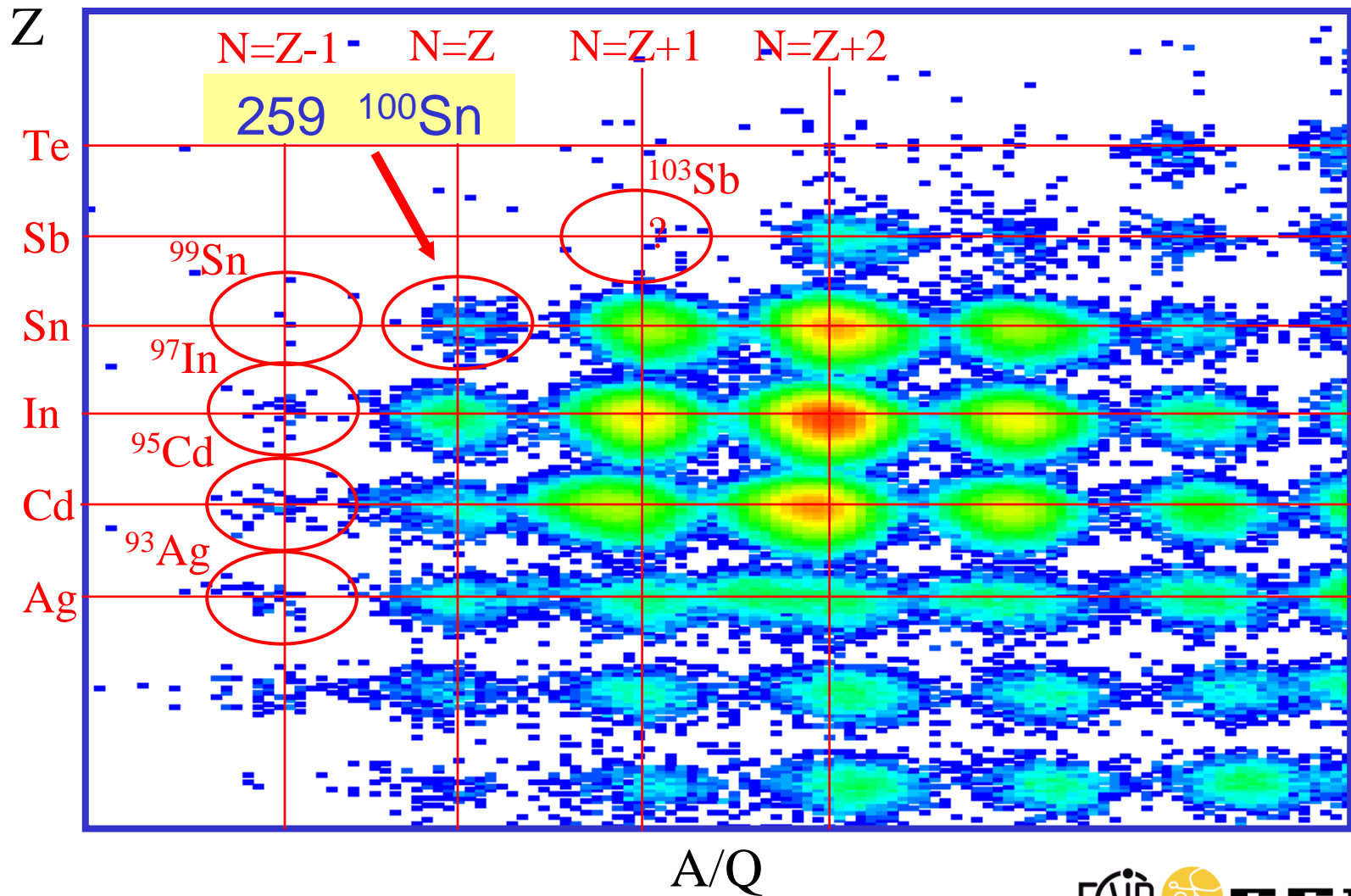
Where does the 1327 keV line come from ???

Active Target

Silicon IMplantation Detector and Beta Absorber

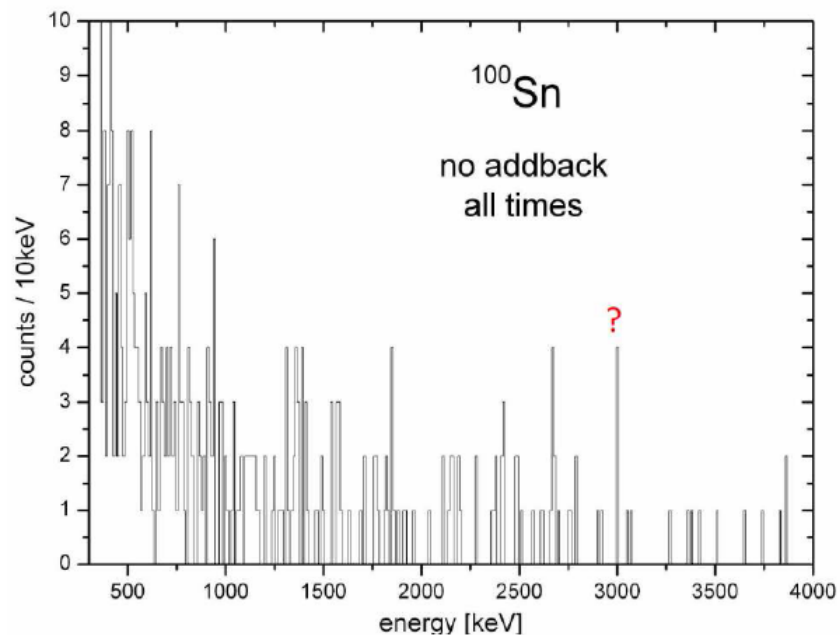
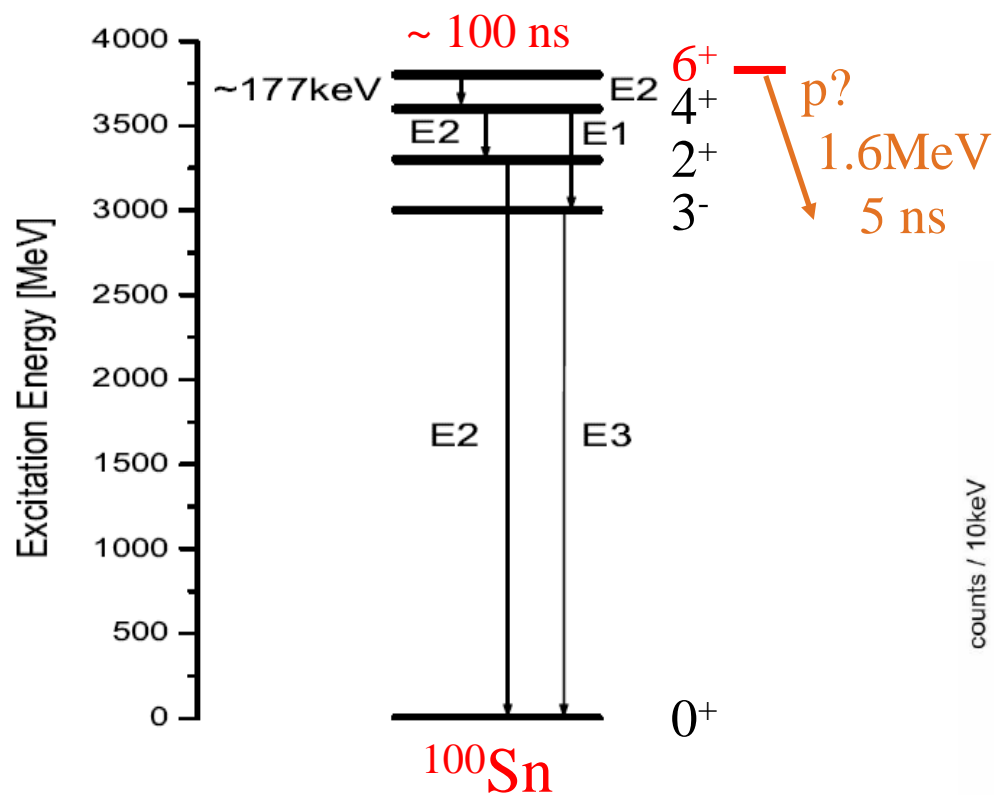


Spectroscopy of the doubly magic nucleus ^{100}Sn and its decay



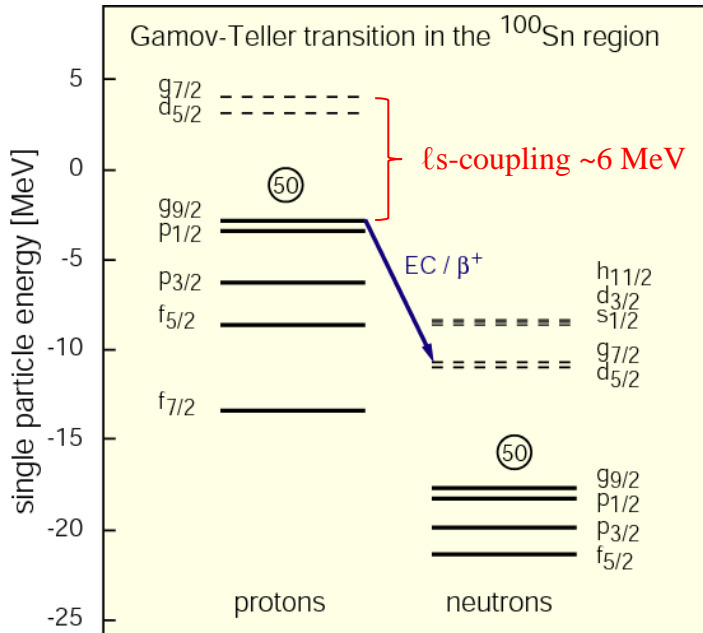
Spectroscopy of the doubly magic nucleus ^{100}Sn and its decay

Theoretical predictions for the ^{100}Sn level schemes



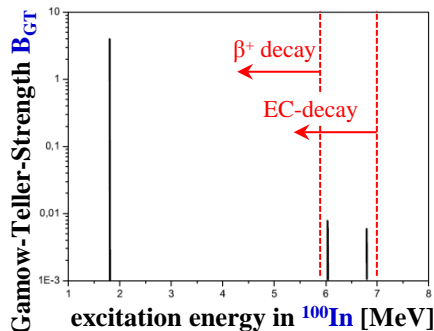
Gamov-Teller strength and Q_{EC} value in the β -decay of ^{100}Sn

Single particle energies for shell model orbitals in ^{100}Sn



$$\beta^+: Q = M(Z+1)c^2 - M(Z)c^2 - 2m_e c^2$$

$$EC: Q = M(Z+1)c^2 - M(Z)c^2 - BE(K\text{-electron})$$



❖ ^{100}Sn is an ideal testing ground to investigate GT-strength:

pure GT spin-flip transition: $0^+ \Rightarrow (\pi g_{9/2}^{-1} \nu g_{7/2}) 1^+$

❖ Almost the whole strength of the GT resonance is covered by the energy window of the β^+ -decay

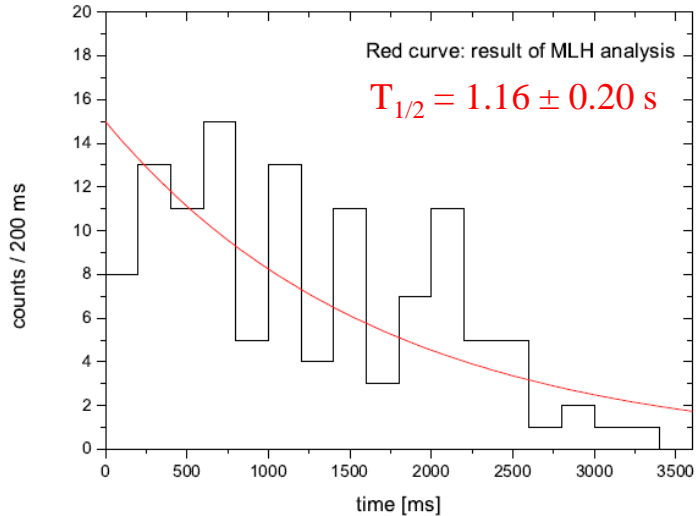
Theoretical calculation of the distribution of the GT-strength:

97% of the whole strength is concentrated in a single state, which is accessible in the β^+ -decay

$$B_{GT}(ESM) = \frac{4\ell}{2\ell + 1} \cdot \left(1 - \frac{N_{\nu g_{7/2}}}{8}\right) \cdot N_{\pi g_{9/2}} = 17.78$$

with $\ell=4$ $N_{\nu g_{7/2}}=0$ $N_{\pi g_{9/2}}=10$

Gamov-Teller strength and Q_{EC} value in the β -decay of ^{100}Sn



The **Gamow-Teller Strength B_{GT}** (only one final state populated) can be calculated from the half life $T_{1/2}$ and the Fermi Phasespace Integral $f(Z, E_0)$:

$$f(Z, E_0) \cdot T_{1/2} = \frac{2\pi^3 \hbar^7}{m_e^5 c^4 G_F^2} \cdot \frac{\ln 2}{g_V^2 \cdot |M_F|^2 + g_A^2 \cdot |M_{GT}|^2}$$

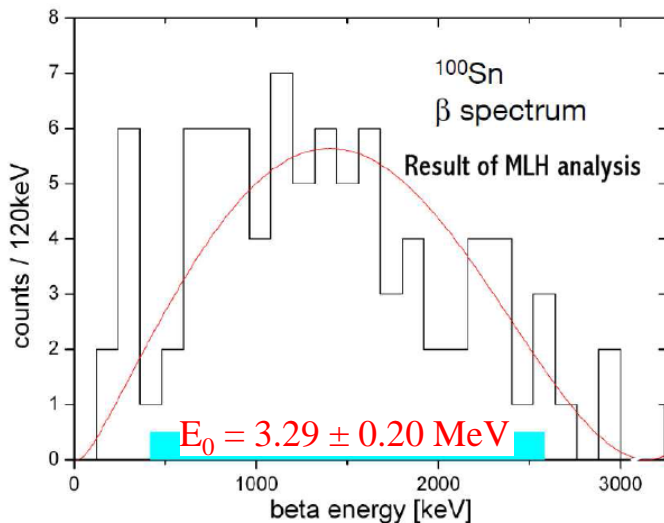
$$G_F/(\hbar c)^3 = 1.16637(1) \cdot 10^{-5} \text{ GeV}^{-2}, \quad g_A/g_V = 1.2695 \pm 0.0029$$

$$f(Z, E_0) \cdot T_{1/2} = \frac{6142.8 \text{ s}}{B_F + (g_A/g_V)^2 \cdot B_{GT}}$$

In the case of a pure Gamow-Teller decay the transition strength can be calculated in the following way:

$$B_{GT} = \frac{3811.5 \text{ s}}{f(Z, E_0) \cdot T_{1/2}} = 9.1^{+4.8}_{-2.3}$$

Fermi-integral with LOGFT program NNDC



Gamov-Teller strength and Q_{EC} value in the β -decay of ^{100}Sn

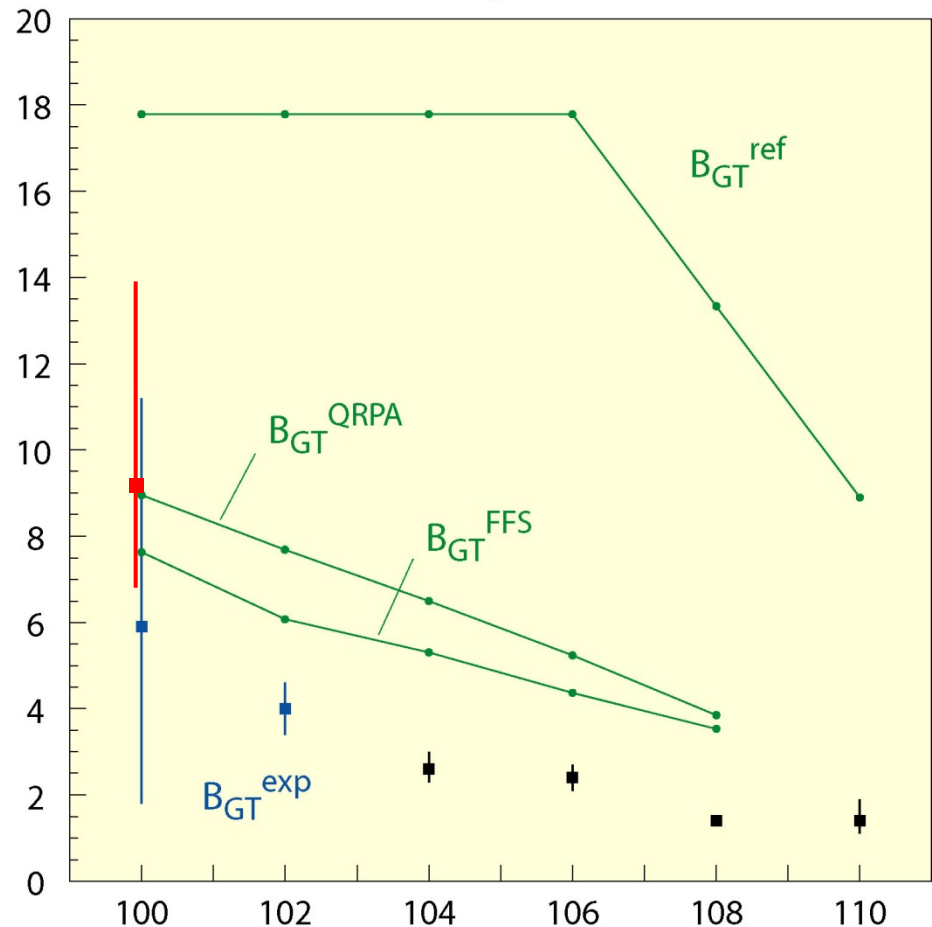
$$B_{GT}(ESM) = \frac{4\ell}{2\ell + 1} \cdot \left(1 - \frac{N_{vg_{7/2}}}{8}\right) \cdot N_{\pi g_{9/2}} = 17.78$$

with $\ell=4$ $N_{vg_{7/2}}=0$ $N_{\pi g_{9/2}}=10$

$$B_{GT} = \frac{3811.5s}{f(Z, E_0) \cdot T_{1/2}} = 9.1^{+4.8}_{-2.3}$$

The main condition for the existence of isolated **Super Gamow-Teller transition** is that the spin-orbit gap between the $\ell+1/2$ and $\ell-1/2$ orbitals (in ^{100}Sn $\ell=4$ orbitals $\pi g_{9/2}, \nu g_{7/2}$) be sufficiently small compared to the shell gap for protons and neutrons (6 MeV), so that the 1-particle-1-hole states are isolated below the 2-particle-2-hole states.

Gamov-Teller strength of Sn isotopes



Gamov-Teller strength and Q_{EC} value in the β -decay of ^{100}Sn

$$Q_{EC} = E_{\beta} + E(1^+) + 2m_e c^2$$

